

In Situ Observations of Solar Wind Stream Interface Evolution

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Abstract The heliocentric orbits of the two STEREO satellites are similar in radius and ecliptic latitude, with separation in longitude increasing by about 45° per year. This arrangement provides a unique opportunity to study the evolution of stream interfaces near 1 AU over time scales of hours to a few days, much less than the period of a Carrington rotation. Assuming nonevolving solar wind sources that corotate with the Sun, we calculated the expected time and longitude of arrival of stream interfaces at the Ahead observatory based

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STEREO Science Results at Solar Minimum

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on the *in situ* solar wind speeds measured at the Behind observatory. We find agreement to within 5° between the expected and actual arrival longitude until the spacecraft are separated by more than 20° in heliocentric inertial longitude. This corresponds to about one day between the measurement times. Much larger deviations, up to 25° in longitude, are observed after 20° separation. Some of the deviations can be explained by a latitude difference between the spacecraft, but other deviations most likely result from evolution of the source region. Both remote and *in situ* measurements show that changes at the source boundary can occur on a time scale much shorter than one solar rotation. In 32 of 41 cases, the interface was observed earlier than expected at STEREO/Ahead.

1. Solar Wind Stream Interfaces

From the earliest *in situ* observations the solar wind was considered to have a bimodal structure; that is, it was divided into two basic populations: slow and fast. Neugebauer and Snyder (1966) first described the average properties of the solar wind based on *in situ* data from *Mariner 2*, launched in 1962. Fast-moving, hot streams (up to 830 km s^{-1} and $9 \times 10^5 \text{ K}$) alternate with slower streams having lower temperatures (down to about 300 km s^{-1} and $3 \times 10^4 \text{ K}$). The fast streams are seen to recur with the Sun's equatorial rotation period of about 27 days. At the leading edge of the fast streams are density enhancements, but outside of these times the density and velocity are inversely related.

It is generally agreed that fast solar wind originates from coronal holes (*e.g.*, Krieger, Timothy, and Roelof, 1973), whereas the slow solar wind comes from the vicinity of the streamer belt. Under the assumption of a dipolar magnetic field aligned with the Sun's rotation axis, a stationary observer in the ecliptic plane will see an alternating pattern of slow and fast wind because the Sun's axis of rotation is tilted approximately 7° with respect to the ecliptic plane.

The transition from slow to fast solar wind is called a "stream interface" (SI), a term coined by Burlaga in 1974. Burlaga (1974) defined a SI as follows: "At 1 AU there is a distinct boundary ... characterized by an abrupt (approximately a factor of 2 change in $<10^6 \text{ km}$) drop in density, a similar increase in temperature, and a small increase in speed." A superposed epoch analysis by Gosling *et al.* (1978) provided further characterization of SIs. They observed that SIs follow 2 to 36 hours after a magnetic field polarity reversal. The SI separates dense, slow-moving plasma from fast, less dense plasma. In the corotating reference frame, there is a velocity shear at the SI, which decreases with radial distance from the Sun as momentum is transferred across the SI. The electron temperature rises by about 40%, and the alpha particle abundance usually increases. Gosling *et al.* (1978) suggested that the fast and slow solar wind streams have different origins. This was supported by later studies such as that of Wimmer-Schweingruber, von Steiger, and Paerli (1997), which showed that solar wind composition changed with the passage of a SI. In addition, Gosling *et al.* (1978) also reported that the plasma in the vicinity of a SI undergoes an east–west deflection in the ecliptic plane. A similar deflection in the out-of-ecliptic (north–south) plane confirmed that SIs are really three-dimensional structures (Gosling *et al.*, 1993).

Although there have been many observations of SIs, multipoint observations are sparse except in the vicinity of L1. Nolte and Roelof (1973a, 1973b) used observations from *Pioneer 8*, *Pioneer 9*, and *Vela* to test the mapping of interplanetary structures back to the Sun, including a case with a high-speed SI. Neugebauer *et al.* (1998) combined measurements from the *Ulysses* and *Wind* spacecraft to determine the latitudinal dependencies of structures in the solar wind and their mapping to the corona using different source surface

models for solar activity minimum time periods. Neugebauer *et al.* (2002) extended this analysis to active times, using data from the *Ulysses* and ACE spacecraft.

During the minimum and rising phases of solar cycle 21, the *Helios 1* and *2* spacecraft provided measurements from 0.3 to 1 AU. These data were combined with observations made near 1 AU. Schwenn (1990) reports that SIs have Parker-spiral-like geometry. Schwenn fit a single Archimedean (Parker) spiral through two observation points and used this to calculate an effective *radial* SI propagation speed as

$$V_{\text{eff}} = \Omega_{\text{Sun}}(R_2 - R_1)/(\varphi_2 - \varphi_1), \quad (1)$$

where Ω_{Sun} is the Sun's angular speed, R_1 and R_2 are the orbital radii of the two observation points, and φ_1 and φ_2 are the longitudes of the observers in the corotating Carrington frame. If the two observation points are not separated in either radius or longitude, then the effective speed is undefined. The separation in Carrington longitude between the STEREO observatories increases with time for the period studied here, but the radial separation waxes and wanes between 0.03 and 0.13 AU.

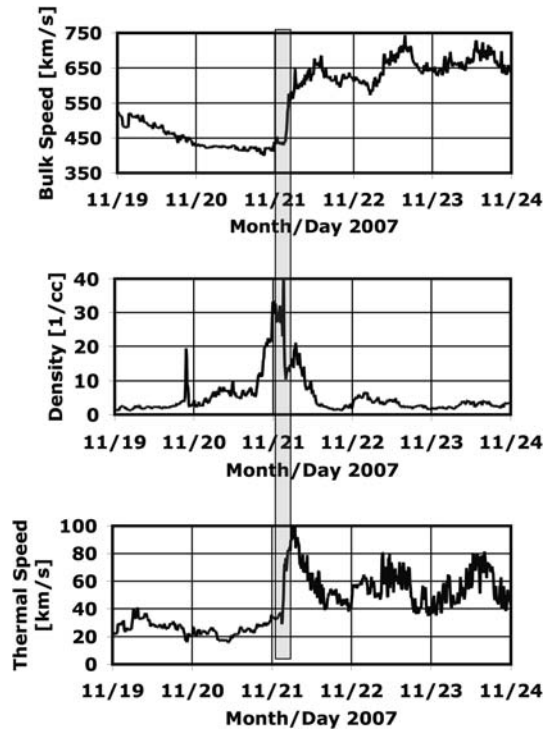
With several dozen interaction regions observed by both STEREO spacecraft at solar minimum, this study will test the scale under which basic assumptions of Parker-spiral geometry are valid near 1 AU. The two primary assumptions are ideal corotation of the SI with the Sun and negligible source evolution. If these assumptions hold it should be possible to predict the arrival time and location of a SI at one observatory after it has been observed at the other by using the mathematical definition of an Archimedean spiral. In addition, the effective propagation speed as calculated by Schwenn (1990) is expected to fall within the range of observed solar wind speeds at 1 AU.

2. The STEREO/PLASTIC Proton Data Set

The *Solar Terrestrial Relations Observatory* (STEREO) was successfully launched in October 2006. STEREO is composed of two nearly identical spacecraft in heliocentric orbit. STEREO/Ahead (STEREO/A) has an orbit just inside 1 AU, and STEREO/Behind (STEREO/B) is just outside 1 AU. The different radii result in orbital periods shorter than one year for STEREO/A and longer than one year for STEREO/B. The heliographic longitude separation between the two observatories increases by about 45° per year. For a more detailed description of the STEREO mission see Kaiser *et al.* (2008). STEREO was launched near the minimum of the solar activity cycle, so the *in situ* observations of the solar wind have been dominated by recurring patterns of slow and fast solar wind with very few transients. This has been a nearly ideal circumstance to study the transitions between slow and fast solar wind streams. The data used in this study are primarily proton data from the Plasma and Suprathermal Ion Composition (PLASTIC) investigation, described in detail by Galvin *et al.* (2008).

The proton parameters are calculated by using a one-dimensional velocity distribution function created from a single coincidence rate on a minute-by-minute basis. A Maxwellian distribution function is fit to the measured distribution function to obtain density, bulk speed, and 1-D temperature for the solar wind protons. The data are then averaged over 10-minute intervals. These bulk proton parameters are used to identify SIs, as described in the next section.

Figure 1 Proton bulk parameters from PLASTIC/A. Top panel is bulk proton speed, middle panel is proton density, and bottom panel is proton thermal speed. Data have been 10 minute averaged. The SI is indicated with shading.



3. Observations

Figure 1 shows a sample SI observed by PLASTIC/A. The SI was identified following the criteria set by Burlaga (1974): an increase in speed accompanied by a drop in density and an increase in temperature. Using these sharp parameter variations we are able to time the SI crossing to within one hour. We have identified in this manner 41 SIs that were observed by both spacecraft between March 2007 and February 2008 (Carrington Rotations 2054 through 2066). Table 1 lists the SI arrival times at STEREO/B, the solar wind speed measured at that time by PLASTIC/B, the time difference between arrivals at the two observatories (with time differences less than zero indicating that STEREO/A was the first observer), the latitude and longitude separations of the observatories in Heliocentric Inertial (HCI) coordinates, the difference in orbital radii in AU, the difference between the expected and actual longitude of arrival at STEREO/A, and the difference between the expected and actual arrival time at STEREO/A. In the HCI coordinate system z is the solar rotation axis, and x is the Sun's ascending node in the ecliptic plane. In other words, the HCI coordinate system does not corotate with the Sun. Latitude differences less than zero indicate that STEREO/A is south of STEREO/B.

The expected longitude of SI arrival at STEREO/A is calculated by assuming the SI is an ideal Archimedean spiral that is propagating at the solar wind speed measured by PLASTIC/B. Any effects from the difference in latitude between the observatories are neglected. The calculation is made in the Carrington coordinate system, using the photospheric equatorial angular speed $\Omega_{\text{Sun}} = 14.38^\circ$ s per day (Newton and Nunn, 1951):

$$\varphi_{\text{Aexpected}} = \varphi_{\text{B}} - \Omega_{\text{Sun}}(R_{\text{B}} - R_{\text{A}})/V_{\text{BSW}}. \quad (2)$$

Table 1 Summary list of stream interfaces.

No.	Interface arrival time at B (UT)	V_{sw} at B (km s^{-1})	Time from B to A (hours)	ΔHCI latitude (degrees)	ΔHCI longitude (degrees)	ΔR (AU)	Expected – observed Carrington longitude at A (degrees)	Expected – observed arrival time at A (hours)
1	5 March 2007 02:05	370	−2.58	−0.07	1.4	0.039	−0.2	−0.4
2	6 March 2007 11:00	445	−3.25	−0.06	1.5	0.040	−1.0	−1.9
3	11 March 2007 23:45	385	−3.92	0.00	1.7	0.043	−1.1	−1.9
4	12 March 2007 14:40	500	−1.00	0.01	1.8	0.044	−0.1	−0.3
5	25 March 2007 08:15	360	−8.08	0.18	2.5	0.053	−3.3	−6.0
6	1 April 2007 04:00	400	−1.42	0.29	3.1	0.058	−0.2	−0.3
7	9 April 2007 10:15	310	−2.58	0.43	3.8	0.065	0.1	0.1
8	22 April 2007 18:40	365	−2.75	0.72	5.1	0.076	−1.4	−2.6
9	23 April 2007 05:15	380	−3.00	0.73	5.1	0.076	−1.2	−3.2
10	27 April 2007 16:00	500	0.83	0.83	5.6	0.080	−1.2	−2.2
11	7 May 2007 15:15	425	−0.50	1.08	6.9	0.088	−2.0	−3.7
12	18 May 2007 07:10	380	7.50	1.36	8.3	0.096	2.1	4.0
13	2 June 2007 17:30	405	1.33	1.75	10.7	0.107	−3.4	−6.3
14	22 June 2007 02:00	445	3.00	2.12	14.2	0.118	−5.9	−10.8
15	29 June 2007 14:30	420	16.75	2.20	15.6	0.122	0.8	1.5
16	3 July 2007 18:25	530	18.33	2.23	16.4	0.123	−0.6	−1.1
17	10 July 2007 20:45	390	14.25	2.23	17.8	0.126	−2.0	−3.7
18	14 July 2007 11:10	460	18.33	2.22	18.6	0.127	−1.7	−3.1
19	20 July 2007 04:45	375	21.42	2.16	19.8	0.128	0.4	0.8
20	26 July 2007 13:25	385	15.67	2.05	21.1	0.129	−4.2	−7.7

Table 1 (*Continued*)

No.	Interface arrival time at B (UT)	V_{sw} at B (km s^{-1})	Time from B to A (hours)	ΔHCI latitude (degrees)	ΔHCI longitude (degrees)	ΔR (AU)	Expected – observed Carrington longitude at A (degrees)	Expected – observed arrival time at A (hours)
21	29 July 2007 07:15	450	4.58	1.99	21.7	0.129	–12.0	–22.0
22	6 August 2007 04:40	330	26.83	1.77	23.3	0.129	1.1	2.0
23	15 August 2007 14:10	470	18.83	1.40	25.3	0.128	–8.2	14.9
24	24 August 2007 23:00	410	44.17	0.93	27.2	0.126	4.0	8.4
25	31 August 2007 18:10	350	26.83	0.52	28.6	0.124	–5.1	–9.2
26	14 September 2007 10:50	370	19.50	–0.41	31.2	0.119	–12.6	–23.3
27	20 September 2007 00:25	435	45.92	–0.83	32.2	0.116	–0.4	–0.8
28	26 September 2007 19:30	430	48.92	–1.36	33.4	0.112	–0.1	–0.2
29	3 October 2007 03:30	540	29.17	–1.87	34.4	0.108	–13.5	–24.6
30	17 October 2007 08:30	440	47.67	–3.00	36.6	0.099	–4.9	–8.9
31	24 October 2007 21:00	515	34.50	–3.57	37.6	0.093	–14.2	–25.9
32	13 November 2007 10:40	480	23.00	–4.81	39.9	0.078	–23.2	–42.4
33	20 November 2007 06:50	410	20.92	–5.14	40.6	0.073	–24.7	–45.1
34	24 November 2007 06:50	510	41.67	–5.30	40.9	0.070	–14.7	–26.7
35	9 December 2007 20:00	530	68.33	–5.66	42.3	0.058	–2.1	–3.8
36	16 December 2007 11:00	490	60.50	–5.68	42.7	0.054	–6.8	–12.4
37	18 December 2007 11:40	520	59.00	–5.67	42.9	0.052	–8.0	–14.7
38	3 January 2008 18:40	450	78.50	–5.30	43.9	0.043	1.5	2.7

Table 1 (Continued)

No.	Interface arrival time at B (UT)	V_{sw} at B (km s^{-1})	Time from B to A (hours)	ΔHCI latitude (degrees)	ΔHCI longitude (degrees)	ΔR (AU)	Expected – observed Carrington longitude at A (degrees)	Expected – observed arrival time at A (hours)
39	10 January 2008 16:10	435	85.33	−5.00	44.3	0.040	4.8	8.7
40	30 January 2008 02:20	360	80.83	−3.72	45.2	0.035	1.4	2.6
41	9 February 2008 10:00	465	57.50	−2.82	45.7	0.034	−12.4	−22.6

The time corresponding to the expected arrival longitude is then obtained by using the STEREO orbit tool (available at <http://stereo-ssc.nascom.nasa.gov/where/>).

From March 2007 through February 2008 the HCI longitude separation between the STEREO observatories grew from less than 2° to just over 48° . The separation in HCI latitude varied but never exceeded 6° . The difference in orbital radii varied from 0.03 to 0.13 AU. The temporal separation between SI observations ranged from 10 minutes to 3.5 days, allowing us to study the evolution of SIs over periods much shorter than a complete solar rotation.

To illustrate the differences in the *in situ* data between the two spacecraft, Figure 2(a) shows the proton bulk speed from STEREO/A and STEREO/B for Carrington Rotations 2061 through 2064, covering approximately mid-September through December 2007. The identified SIs for the two spacecraft are marked with vertical lines (red for STEREO/A and blue for STEREO/B) and labeled with the numbers corresponding to their listings in Table 1. The longitudinal separation between STEREO/A and STEREO/B during this time increases from 33° to 45° . Assuming a rigid corotation rate of 14.38° per day, and a typical garden hose angle of 45° at 1 AU, we would expect the time difference between observations is to increase from about 45 to about 73 hours. (These values are less than the longitude difference divided by the Sun's angular speed because of the small difference in orbital radii between the observatories.) The actual time differences ranged from 21 hours (SI no. 33) to 68 hours (SI no. 35).

To compare the data between observatories in more detail, we have plotted proton bulk speed versus back-mapped Carrington longitude in Figure 2(b). In the “ballistic” back-mapping technique it is assumed the solar wind has traveled with constant speed radially outward from the solar source surface, as described by Nolte and Roelof (1973a). The orbital radius of the spacecraft is divided by the measured solar wind speed to obtain a travel time. The Sun's equatorial angular speed is divided by the travel time to obtain the difference in longitude between the source and the observatory. This method of estimating the source longitude has been used in previous studies including those by Nolte and Roelof (1973a, 1973b) and Neugebauer *et al.* (1998, 2002). The error in the source longitude calculated using this technique is estimated to be $\leq 10^\circ$ (Nolte and Roelof, 1973a). Note that in Figure 2(b) the back-mapped slow-to-fast solar wind transitions are stretched across a wide range of longitudes compared to the real-time *in situ* measurements in Figure 2(a). This is not unexpected; material from a range of longitudes at the source is compressed as the fast solar wind tries to overtake the slow wind in corotating interaction regions (CIRs).

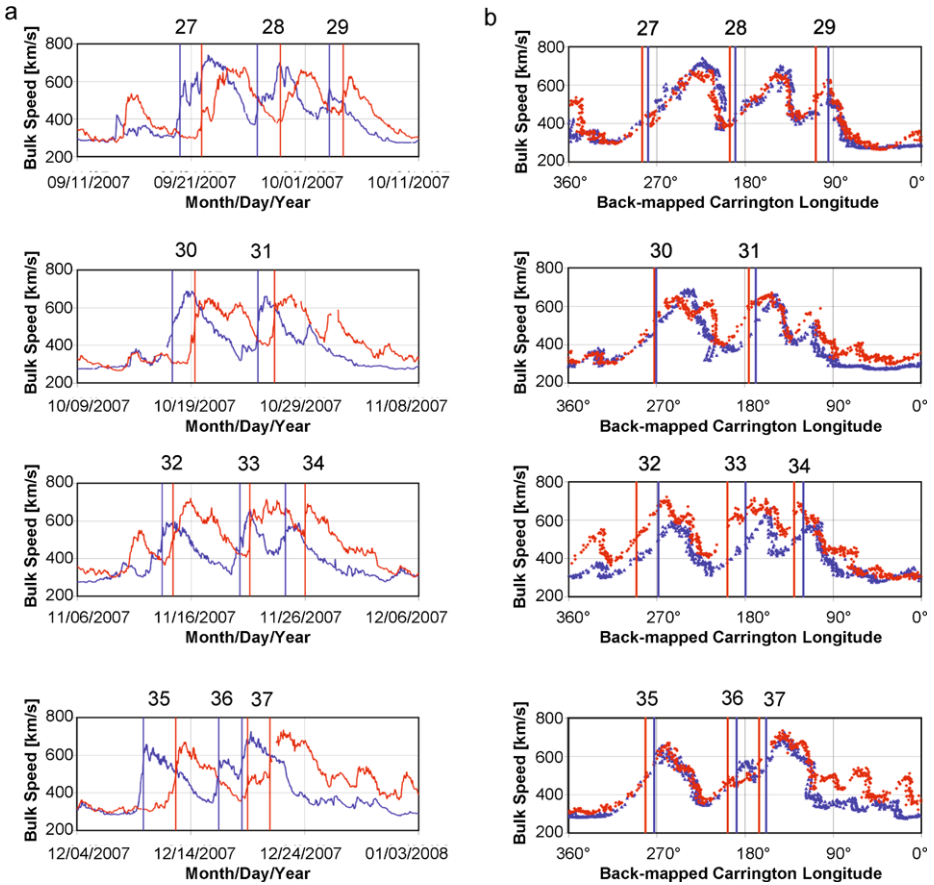


Figure 2 Proton bulk speed versus time (a) and back-mapped Carrington longitude (b) for Carrington Rotations 2061 through 2064. The red trace is STEREO/A, and the blue trace is STEREO/B. Vertical bars indicate the positions of SIs. Note that the back-mapped Carrington longitudes are reversed from typical synoptic maps.

Under the assumptions of perfect corotation and negligible source evolution, the back-mapped speed profiles in Figure 2(b) are expected to match between the two observatories. Recurring streams (and their associated CIRs) should be observed at nearly the same Carrington longitudes for successive solar rotations under the same assumptions. Vertical lines in Figure 2(b) again indicate the SIs identified for this study. Note that although only the solar wind speed is shown, there are additional density and temperature criteria for SI identification. As a result, even when speeds are the same for both spacecraft, the back-mapped longitude of the SIs may not agree, such as for SI no. 35.

In this back-mapped representation, the large- and small-scale features of each Carrington rotation can be compared. In Figure 2 each Carrington rotation begins and ends with slow solar wind. The middle of each rotation is dominated by fast solar wind with three high-speed peaks and short periods of slower wind between them. The high-speed peaks move slightly from right to left with each Carrington rotation, indicating they recur sooner than expected (again under the assumption that they corotate with the Sun's equatorial an-

gular speed of 14.38° per day). This early recurrence could be explained by growth of the source and/or by movement of the source with respect to the photosphere.

In the top panel of Figure 2(b), Carrington Rotation (CR) 2061, the first two high-speed streams (corresponding to SI nos. 27 and 28) show very similar speed profiles for STEREO/A and STEREO/B. The back-mapped longitudinal extent is the same for both observatories. The maximum observed solar wind speed is about 60 km s^{-1} greater at STEREO/B than at STEREO/A. The back-mapped longitudes of the SIs agree within 6° between the two observatories. The third SI, no. 29, shows more differences between the two observatories. SI 29 back-maps to 108° Carrington longitude from STEREO/A and to 95° Carrington longitude from STEREO/B. The high-speed stream following SI 29 back-maps to a region covering 10° more longitude from PLASTIC/A than from PLASTIC/B. The maximum observed solar wind speed is about 630 km s^{-1} at A and 560 km s^{-1} at B. The observatories were separated by about 33° in HCI longitude and by about 2° in HCI latitude.

In CR 2062, the first SI (no. 30) maps to nearly identical longitudes (271° and 272°) from the two spacecraft, but the speed profiles following the SI are obviously different. PLASTIC/A records two high-speed peaks, but PLASTIC/B saw just one. The second transition (SI no. 31) maps back to about 176° from A and to 169° from B. The high-speed stream is about 10° wider in back-mapped longitude from STEREO/A than from STEREO/B. The third transition from slow to fast solar wind is not marked with a SI, nor is it given a number in Table 1 because the SI criteria of a simultaneous drop in density and increase in temperature accompanying an increase in speed were not met.

The Carrington rotation in Figure 2 with the greatest apparent discrepancies in speed profiles and SI longitudes between Ahead and Behind is CR 2063 (November 2007), when HCI longitudinal separation was about 40° and HCI latitude separation was about 5° . By CR 2064 (bottom panel of Figure 2) the speed profiles are showing better agreement again. The longitudinal separation between the observatories for CR 2064 increased to about 43° , and the latitudinal separation was about 5° . Possible explanations for the different speed profiles in CR 2063 include the following: movement of the source (*i.e.*, nonideal corotation), evolution of the source (such as shrinking or growing surface area), which can result in a change in the radial speed at which material leaves the source (*cf.* Nolte *et al.*, 1976; Wang and Sheeley, 1990), and back-mapping to different source latitudes at the Sun. According to the STEREO/IMPACT level 3 event list (available at http://www-ssc.igpp.ucla.edu/forms/stereo/stereo_level_3.html), no interplanetary coronal mass ejections were observed during CR 2063, so the possibility of transient events will not be explored.

We begin by examining the sources of the fast streams: the coronal holes. Some coronal holes rotate rigidly, independent of the photosphere (see, for example, Timothy, Krieger, and Vaiana, 1975). Others have latitude-dependent differential rotation similar to the photosphere (see, for example, Wang and Sheeley, 1993). Nolte *et al.* (1976) showed that the speed of the solar wind emanating from a coronal hole is proportional to the area of the hole within 10° of the ecliptic plane. Similarly, Wang and Sheeley (1990) showed that the total coronal hole area is correlated with observed solar wind speeds at 1 AU. Figure 3 shows synoptic plots of modeled coronal holes obtained from the Global Oscillation Network Group (GONG) for the same Carrington rotations as Figure 2. To create these plots, a potential-field source surface model with the source surface fixed at 2.5 solar radii (R_{Sun}) is determined from hourly synoptic solar magnetograms. Purple and orange shading represent regions of open fields (*i.e.*, coronal holes) with opposite magnetic polarity. The solid black curve is the modeled location of the heliospheric current sheet (HCS). Green lines indicate

fields that are closed out to just less than $2.5 R_{\text{Sun}}$. In terms of Carrington longitude, observation time goes from right to left. In other words, a feature observed at CR 2063.25 would be found at 270° Carrington longitude on the synoptic map for CR 2063. In each synoptic map there is an equatorial coronal hole extending from the northern pole in the vicinity of longitudes 225° to 270° . There is a broader coronal hole extending from the southern pole between about 80° and 180° longitude. Both regions are shaded in Figure 3. Note that although the large-scale features are the same, the boundaries of the holes do change from one solar rotation to the next. The equatorward extent of the modeled southern coronal hole moves to the right (west), crossing the edge of the shaded region in CR 2064. Similarly, the equatorward extent of the modeled northern hole also moves to the right between CR 2061 and 2062.

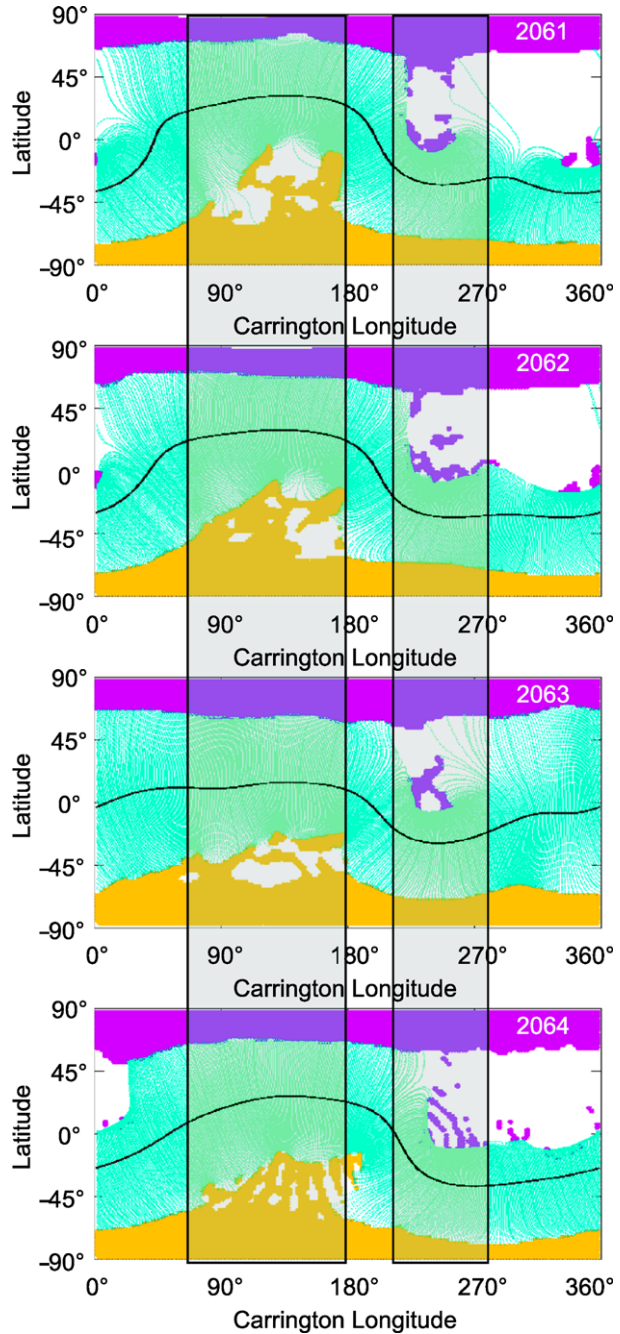
Figures 4, 5, 6 and 7 show the same GONG modeled coronal hole maps, combined with plots of back-mapped solar wind speed measured by PLASTIC. The *in situ* data are plotted in time from right to left for easier comparison with the synoptic maps. The middle two panels of Figures 4 through 7 show the back-mapped radial magnetic field polarity measured by STEREO/IMPACT (see Luhmann *et al.*, 2008). The bottom two panels are synoptic maps created from 195 Å wavelength extreme ultraviolet (EUV) observations from STEREO/SECCHI (see Howard *et al.*, 2008), with Ahead images above and Behind images below. Regions of open magnetic field in the EUV images appear darker than closed fields, indicating the locations of coronal holes. The images from both STEREO/A and STEREO/B can be compared with the modeled coronal holes on the GONG synoptic maps. The synoptic EUV images can also be compared with each other to look for changes that have taken place during the interval between observations at the two observatories. Comparing images between the two observatories should be done with caution; the apparent size of the disk and the viewing angles are different between the two observatories. Also note that the time at which the image at a particular longitude is taken does not always correspond to the back-mapped time the material measured *in situ* left the source region. The nearly horizontal tracks across the GONG maps in Figures 4 through 7 are projections of the spacecraft trajectories. Vertical lines again indicate the back-mapped longitudes of the SIs, and the numbers correspond to Table 1.

In Figures 4 through 7 the fast solar wind comes from the longitudes corresponding to the coronal holes, as expected. The right-most velocity peak corresponds to the northern coronal hole. The middle peak corresponds to the first encounter with the southern coronal hole. There is a velocity minimum, followed by the left-most peak corresponding to a second encounter with the southern coronal hole. We will concentrate our analysis on the middle peaks (SI nos. 28, 31, 33, and 36), as they show clear SIs between slow and fast solar wind. We will first examine whether latitude dependence can explain the observed changes. We will then look in more detail at source motion and evolution.

If the SI at the Sun is not vertical (*i.e.*, oriented north–south), then the slope of the SI combined with the latitudinal separation between the observatories can result in an apparent longitudinal offset in the observations. Leske *et al.* (2008) and Mason *et al.* (2009) note the importance of latitude separation when comparing CIR energetic particle data between the STEREO observatories. This idea is illustrated in Figure 8 and was described before by Lee (2000). Because the SI slope is not remotely observed, a potential proxy is the slope of the HCS at the longitude where it was crossed. (The GONG consortium models the neutral line at the source surface.) As Gosling *et al.* (1978) noted, SIs are generally observed several hours after a current sheet crossing. We tried this technique to correct for latitude separation with mixed results.

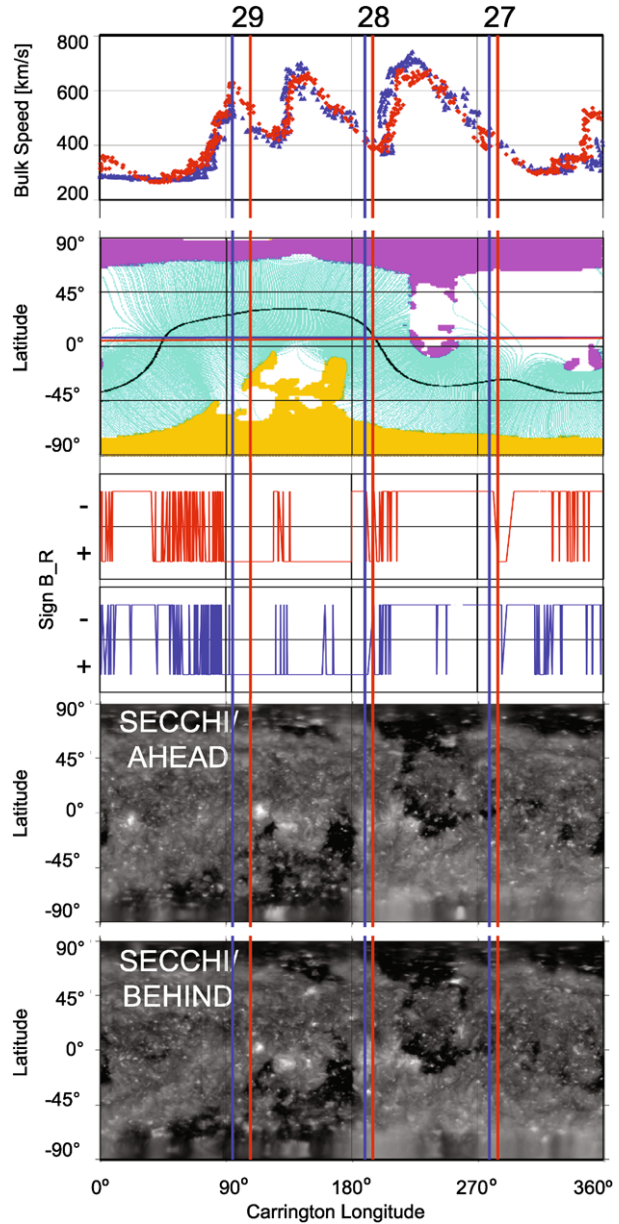
Both STEREO observatories crossed the current sheet just before encountering SI no. 28 in the middle of CR 2061. The slope of the line tangent to the modeled current sheet is about

Figure 3 GONG synoptic coronal hole maps for Carrington Rotations 2061 through 2064. The orange and purple shading denote opposite polarities (orange for positive and purple for negative). The green lines are magnetic field lines that are closed at the model's source surface. White areas represent field lines that close below the source surface radius of $2.5R_{\text{Sun}}$. The solid black curve is the heliospheric current sheet.



-1.80. The observatories were separated by 1.5° HCI latitude at this time. The extra back-mapped longitude ($\Delta\varphi$) that would have to be covered by spacecraft B before encountering

Figure 4 The top panel is the observed solar wind speed back-mapped to its source longitude. The second panel is the GONG synoptic plot of modeled coronal holes and the modeled neutral line for Carrington Rotation 2061. The horizontal red and blue traces are the projections of the STEREO spacecraft trajectories. The middle panels are the sign of the radial component of the magnetic field observed by STEREO/IMPACT and mapped back to the source longitude in the same manner as the proton bulk speed. The bottom two panels are 195 Å SECCHI observations, with SECCHI-A above and SECCHI-B below. Vertical lines indicate the back-mapped longitudes of the SIs, with numbers corresponding to Table 1. Note that time goes from right to left in all of the plots.



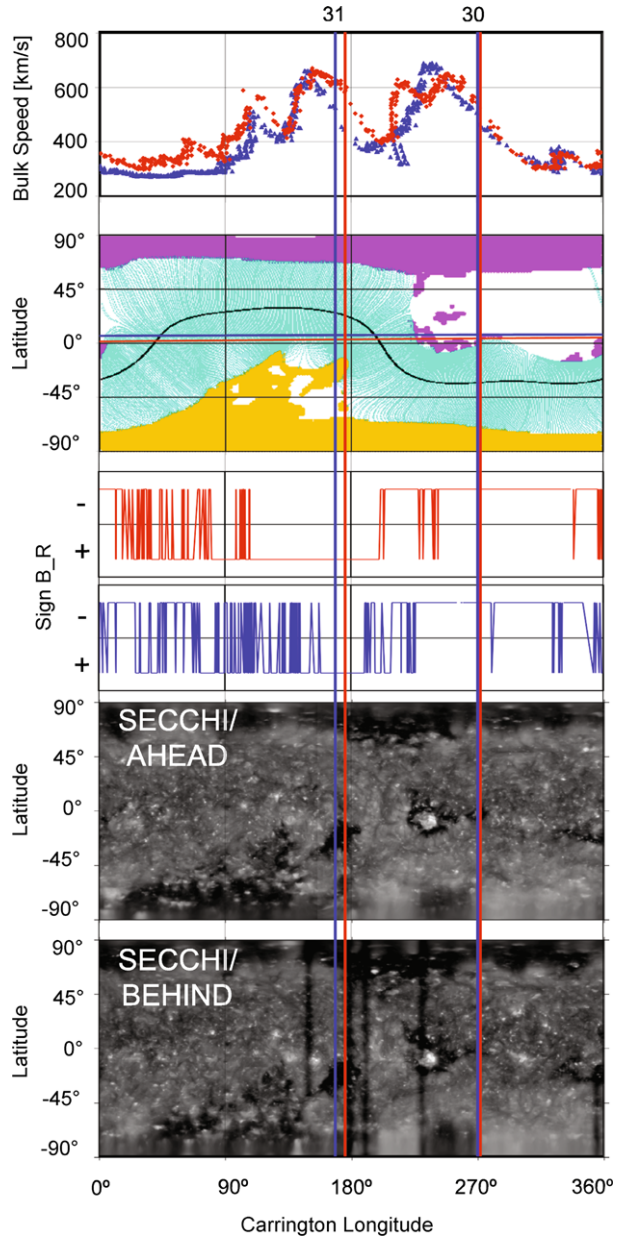
the SI is 7.3° is calculated following Lee (2000) as

$$\Delta\varphi = \frac{\Omega_{\text{Sun}}(R_B - R_A)}{V_{\text{sw}}} + \frac{\theta_A - \theta_B}{\text{slope}}. \tag{3}$$

The actual offset in back-mapped Carrington longitude is 5.6° .

During CR 2062 the observatories crossed the current sheet at about 200° back-mapped Carrington longitude, just before encountering SI no. 31. The slope of the modeled neutral

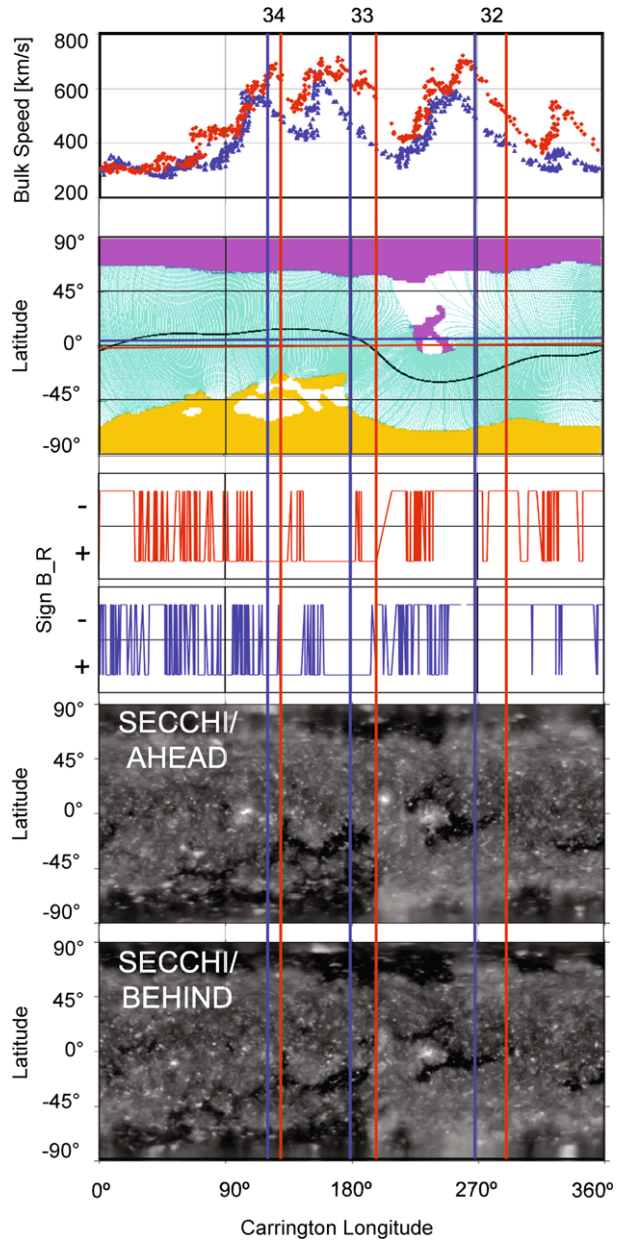
Figure 5 Same as Figure 4, but for Carrington Rotation 2062.



line is approximately -1.68 . The expected offset in longitude is 6.8° . The actual offset in back-mapped Carrington longitude is 7.1° , which agrees well with the calculated value. In this case using the slope of the modeled neutral line (in place of a real SI slope) seems to work.

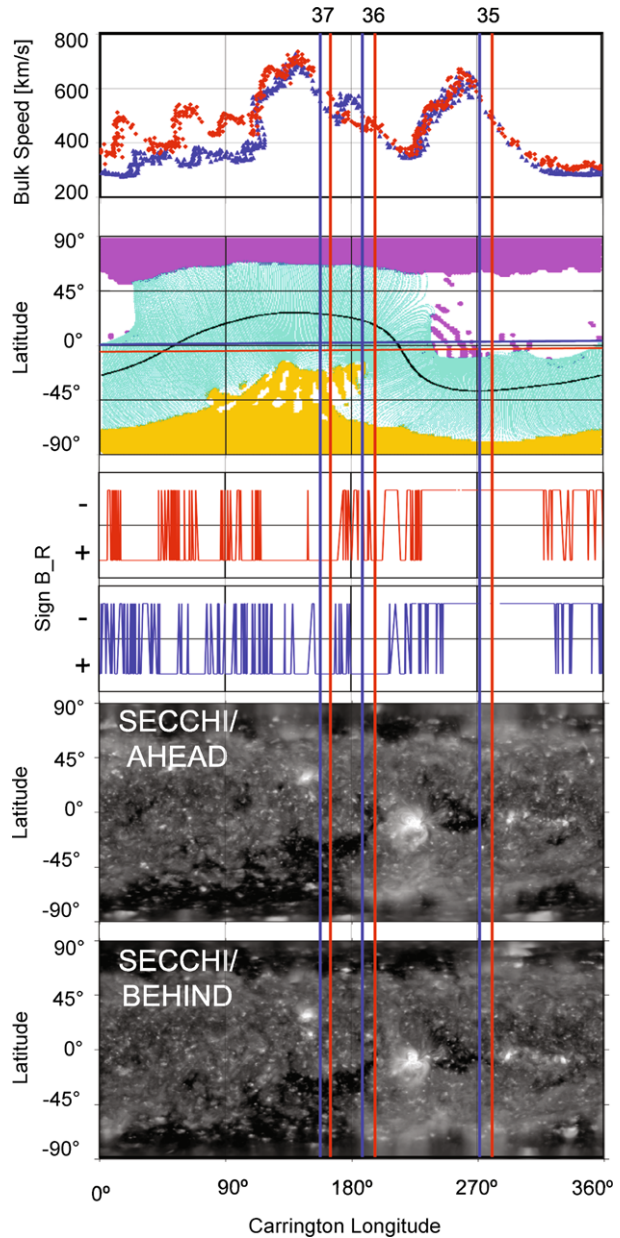
One Carrington rotation later (2063) the slope of the modeled current sheet has changed such that it is less steep, about -1.00 . The HCI latitude separation between the two spacecraft has increased to about 5.25° . The same calculation as before yields an expected longi-

Figure 6 Same as Figure 4, but for Carrington Rotation 2063.



tude offset of 9.6° . The actual offset in back-mapped Carrington longitude for SI no. 33 is about 18.4° , almost twice the calculated value. The SI slope would need to be 0.37 to obtain the larger offset, which is not an unreasonable value. However, from CR 2061 through 2064 the slope of the modeled neutral line only varies from -1.0 to -1.82 . Either latitude separation is not the only cause for the offset in back-mapped longitude or the neutral line slope is not a good substitute for the SI slope in this case.

Figure 7 Same as Figure 4, but for Carrington Rotation 2064.



Finally, the modeled current sheet prior to SI no. 36 in the middle of CR 2064 has a slope of about -1.82 , similar to the slope during CR 2062, but steeper than in CR 2063. The expected offset in back-mapped Carrington longitude is 6.0° . The actual offset is 9.0° .

Another possible reason for the timing difference between the spacecraft (also related to latitude separation) is the latitude of the current sheet. A spacecraft on the positive-polarity side of the current sheet would not see the high-speed flow from a coronal hole with negative polarity. In the case of SI no. 33, the current sheet is so near to the ecliptic plane that

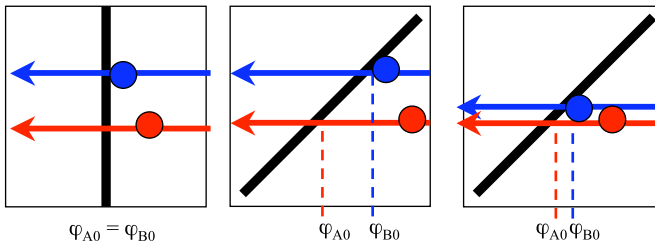


Figure 8 Simple illustration to demonstrate the importance of the slope of the SI when the two observatories are separated in latitude. Each panel shows the out-of-ecliptic plane orthogonal to the Sun – observer line. The SI is the solid black line. The projected spacecraft trajectories are the red and blue arrows.

the GONG model shows it may fall between the spacecraft from Carrington longitudes of 180° to 190° (see the second panel of Figure 6). However, plotting the radial magnetic field polarity recorded by IMPACT (the middle two panels of Figure 6), we see that the observed radial magnetic field polarity at both spacecraft reverses at about 200° . So the solar wind observed by spacecraft B maps to the correct side of the current sheet to observe the coronal hole material, thus excluding current sheet location as a possible explanation for the disparity in back-mapped SI longitude. Note that the right-hand (western) boundary of the southern coronal hole in the SECCHI maps extends to larger Carrington longitudes than the modeled coronal hole in the GONG map, suggesting the model does not agree well with observations. This may explain why the correction for latitude separation did not work in the second example here.

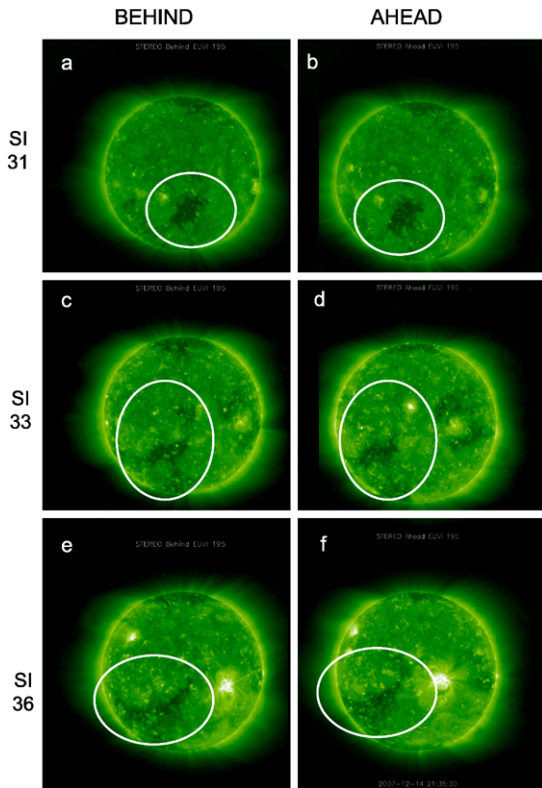
We find that latitudinal differences may explain some of the differences in the *in situ* data between the two spacecraft. However, in some cases, particularly in the case of SI no. 33, the time differences are larger than expected. We now look at how the source is changing during these time periods to determine whether that also plays a role.

In CR 2061 (Figure 4), the velocity profile for the middle peak, corresponding to the southern coronal hole and SI no. 28, is very similar between spacecraft A and spacecraft B. SI no. 28 was back-mapped to 195° Carrington longitude from STEREO/A and to 190° from STEREO/B. The SI was observed shortly after a region of mixed polarity, indicating current sheet material. There is reasonable agreement between the back-mapped current sheet crossing longitude of about 200° Carrington longitude and the predicted longitude from the GONG modeled current sheet. The SECCHI synoptic images show coronal holes with the same locations and boundaries as the GONG model indicates. There are no obvious differences between the SECCHI/A and SECCHI/B images. The HCI latitudinal separation of the two spacecraft in this case was 1.5° .

In CR 2062 (Figure 5), the difference between the STEREO/A and STEREO/B back-mapped longitudes of SI no. 31 (the recurrence of SI no. 28) has increased to about 8° . As mentioned earlier, the high speed observed by STEREO/A back-maps to a larger range of Carrington longitudes than from STEREO/B. The latitude separation between the spacecraft is 3.7° . The top two panels in Figure 9 shows the 195 \AA observations from SECCHI at the times corresponding to the SI departing the source surface (based on ballistic back-mapping) during CR 2062. The images have been scaled so that the apparent disk size is the same for both observatories. A qualitative comparison of the images from STEREO/B and STEREO/A show the boundaries of the coronal hole do not appear to change substantially in the two days between SECCHI/B and SECCHI/A observations.

The top panel in Figure 6 shows that the peak speed observed by PLASTIC/A is larger than the peak speed observed by PLASTIC/B by 50 km s^{-1} following SI no. 33. The Nolte

Figure 9 SECCHI 195 Å images corresponding to the times when the SIs 31, 33, and 36 left the source surface (based on ballistic back-mapping): (a) 21 Oct. 2007 07:46:17 (CR 2062.53 = 169°), (b) 23 Oct. 2007 03:55:30 (CR 2062.51 = 176°), (c) 15 Nov. 2007 21:36:07 (CR 2063.50 = 180°), (d) 17 Nov. 2007 13:25:30 (CR 2063.45 = 198°), (e) 12 Dec. 2007 20:25:58 (CR 2064.48 = 188°), and (f) 14 Dec. 2007 21:35:30 (CR 2064.45 = 198°).



et al. (1976) relationship between maximum solar wind speed and coronal hole area is

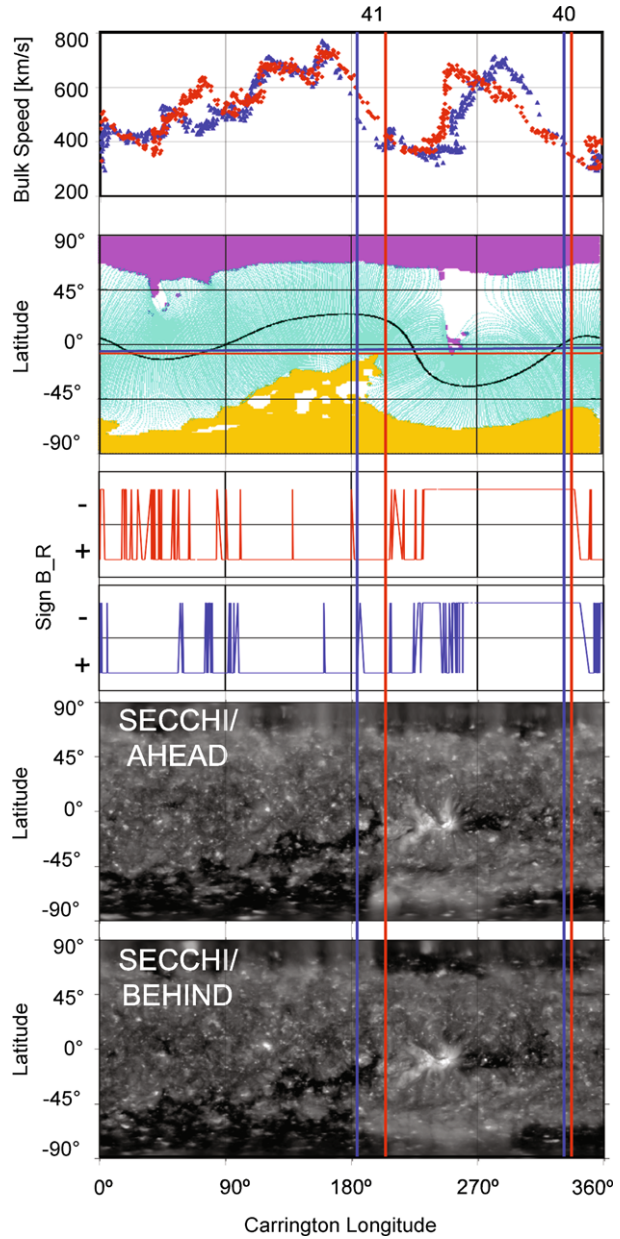
$$V_{SW} = (80 \pm 2)A + 426 \pm 5, \tag{4}$$

where V_{SW} is the maximum solar wind speed in km s^{-1} and A is the area of the coronal hole within 10° of the ecliptic plane in units of 10^{10} km^2 . An increase in maximum solar wind speed of 50 km s^{-1} corresponds to an increase in coronal hole area of $6.25 \times 10^9 \text{ km}^2$, which is about 17% of the Sun's surface area within 10° of the ecliptic plane. The SECCHI 195 Å observations in the middle panels of Figure 9 show that the boundaries of the coronal hole have changed between the observations. In addition to the changing boundaries, there is an active region in the image from SECCHI/A that is not seen in the image from SECCHI/B.

In the subsequent Carrington rotation, 2064 (Figure 7), SI no. 36 (the recurrence of no. 33) maps to about 197° Carrington longitude from both A and B. (This is the same longitude as the mapping from A the previous rotation.) The GONG synoptic map shows that the boundary of the modeled coronal hole has moved to the right (i.e., in the direction of solar rotation) by about 15° . However, the SECCHI synoptic maps in the bottom panel of Figure 7 show that the boundary has not obviously moved from left to right since the previous solar rotation. The SECCHI images in the bottom panels of Figure 9 show qualitative agreement between the two observatories.

Thus we find that our example with the largest time lag discrepancy, SI no. 33, does show a change in the source. However, because it was also a case with a large latitude difference, there is an ambiguity on which is the primary cause. We therefore show one more example to better distinguish between latitudinal dependence and source motion.

Figure 10 Same as Figure 4, but for Carrington Rotation 2066.



During CR 2066 (February 2008) the radial magnetic polarities line up fairly well between the two spacecraft, as shown in the middle panels of Figure 10. The HCI latitude separation between the spacecraft was less than 3°, and the modeled current sheet was not between the observatories. Even so, the profile of the high-speed stream following SI no. 40 (near 305° longitude) is obviously different between the two spacecraft. With such a small difference in latitude, evolution of the source between the two observations seems to be the most likely cause for differences in the speed profiles. The SECCHI observations shown in

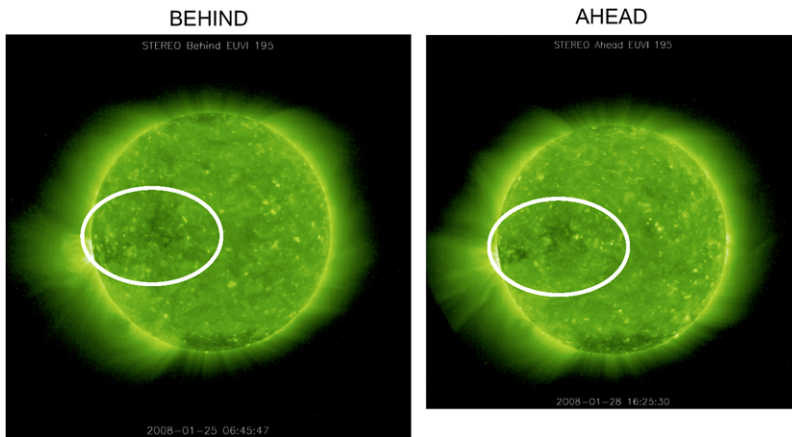


Figure 11 Same as Figure 9, but for SI no. 40.

Figure 11 suggest that the coronal hole is shrinking between the observations by SECCHI/B and SECCHI/A.

4. Discussion

Figure 12 shows the difference between expected and observed SI longitude at STEREO/A plotted versus time. The expected longitude was calculated by ignoring latitude separation between the STEREO observatories (as described in Section 3). Through the end of July 2007, all events except one (no. 14) were measured within 5° of the expected longitude. In July of 2007, the observed time separation is about 20 hours. From that point on, most SIs are observed within 15° of the expected longitude. The two events well outside that range (nos. 32 and 33) were both observed during CR 2063, for which qualitative examination of SECCHI images showed coronal hole evolution taking place (Figure 9). Interestingly, the vast majority of cases were observed at earlier longitudes than expected.

One way to obtain a shorter-than-expected time between observations would be if the SIs were propagating slower than expected, leading to an overwound Parker spiral configuration. The phenomenon of SI overwinding was reported by Nolte and Roelof (1973b). Deviations of the interplanetary magnetic field from ideal Archimedean spiral geometry have been reported by Smith and Bieber (1991). Figure 13 shows the effective SI propagation speeds, calculated using Schwenn's technique (Equation (1)), for each of the 41 SIs listed in Table 1. The effective speeds have been overlaid on a plot of the solar wind speed measured at STEREO/A. We find that through July 2007 there was good agreement between the calculated SI speed and the measured solar wind speed, as expected. However, from August 2007 on, when the longitudinal separation was greater than 22° , and the time separation for SIs was more than a day, we find that the effective propagation speeds of the SIs often fall outside the range of measured solar wind speeds. The effective speeds are in most cases less than the measured solar wind speed, which would be consistent with overwinding and early arrival at STEREO/A. However, if the SI speeds were really significantly less than the solar wind speed, we would observe this independent of the time between the two measurements. In fact, the most accurate measurements of the radial propagation speed are from the time periods when the spacecraft are closest together in longitude. Because we only calculate the

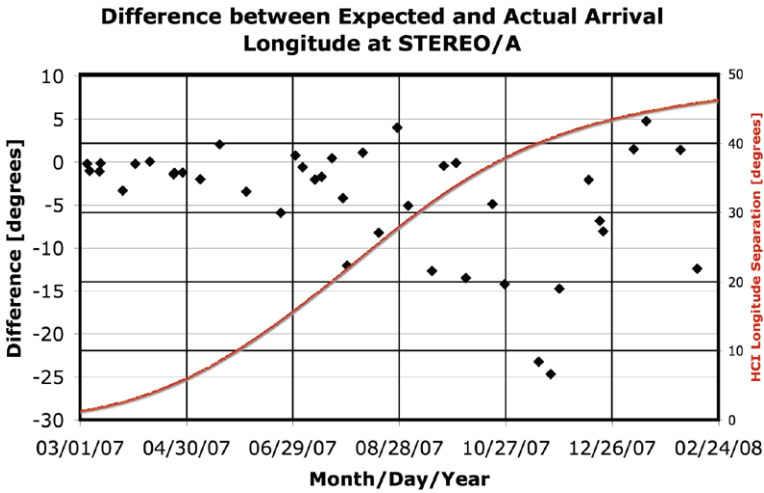


Figure 12 The difference between the calculated and expected observation longitudes versus time, overlaid with a plot of the STEREO HCI longitude separation.

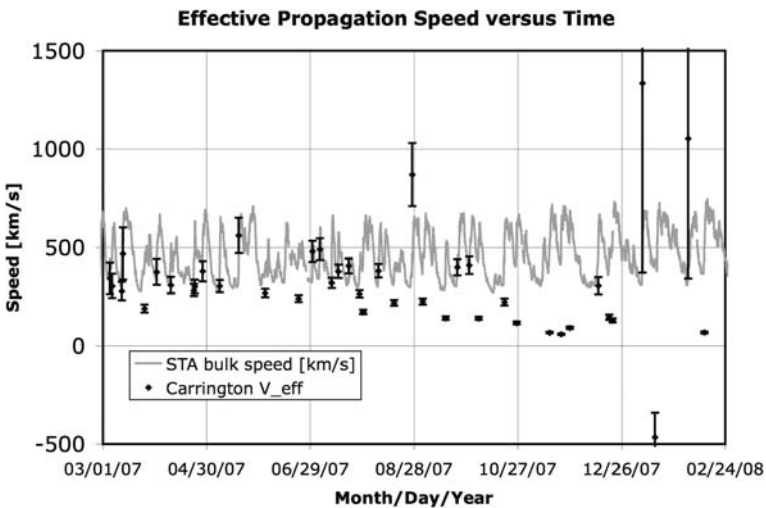


Figure 13 Effective propagation speed versus time. The light gray trace is the solar wind speed measured at STEREO/A. The calculated effective speeds begin to diverge from the measured speeds at the end of July 2007.

low speeds once there is a significant time lag between the two spacecraft measurements; we conclude that this is a result of time or latitude dependence, not a result of a slow SI speed.

5. Summary and Conclusions

Forty-one SIs (listed in Table 1) were observed by both PLASTIC/A and PLASTIC/B from March 2007 through February 2008. The effective propagation speeds of each SI were cal-

culated following the technique of Schwenn (1990). When longitudinal separation between the observatories exceeded about 20° , the calculated effective speeds began to diverge from the measured speeds.

The expected observation longitude at STEREO/A was calculated based on the solar wind speed measured with the SI at STEREO/B. We found that the difference between expected and observed longitudes exceeded 5° for cases when the effective propagation speed was outside the range of measured speeds. In 32 of 41 cases the SI arrived at STEREO/A earlier than expected. This result is important to forecasting the arrival time of compression regions and high-speed streams at Earth, which are known to cause geomagnetic storms.

In some cases the difference between expected and observed longitudes (and differences in observed speed profiles) can be explained by the difference in spacecraft HCI latitude. However, there are cases where the latitude separation is small and the two observatories see obviously different *in situ* solar wind speed profiles and different back-mapped Carrington longitudes for the SI. For the cases with the largest differences, evidence for changes in the coronal holes were observed in the EUV images. Thus, both latitude separation and source evolution can be important in determining when a high-speed stream will reach a particular location.

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