

Competing orders in M-theory: superfluids, stripes and metamagnetism

Aristomenis Donos,^a Jerome P. Gauntlett,^a Julian Sonner^b and Benjamin Withers^c

^a*Blackett Laboratory, Imperial College,
London, SW7 2AZ, U.K.*

^b*C.T.P., Massachusetts Institute of Technology,
Cambridge, MA 02139, U.S.A.*

^c*Centre for Particle Theory and Department of Mathematical Sciences, University of Durham,
South Road, Durham, DH1 3LE, U.K.*

E-mail: a.donos@imperial.ac.uk, j.gauntlett@imperial.ac.uk,
sonner@mit.edu, b.s.withers@durham.ac.uk

ABSTRACT: We analyse the infinite class of $d = 3$ CFTs dual to skew-whiffed $AdS_4 \times SE_7$ solutions of $D = 11$ supergravity at finite temperature and charge density and in the presence of a magnetic field. We construct black hole solutions corresponding to the unbroken phase, and at zero temperature some of these become dyonic domain walls of an Einstein-Maxwell-pseudo-scalar theory interpolating between AdS_4 in the UV and new families of dyonic $AdS_2 \times \mathbb{R}^2$ solutions in the IR. The black holes exhibit both diamagnetic and paramagnetic behaviour. We analyse superfluid and striped instabilities and show that for large enough values of the magnetic field the superfluid instability disappears while the striped instability remains. For larger values of the magnetic field there is also a first-order metamagnetic phase transition and at zero temperature these black hole solutions exhibit hyperscaling violation in the IR with dynamical exponent $z = 3/2$ and $\theta = -2$.

KEYWORDS: Black Holes in String Theory, AdS-CFT Correspondence, Holography and condensed matter physics (AdS/CMT)

ARXIV EPRINT: [1212.0871](https://arxiv.org/abs/1212.0871)

Contents

1	Introduction	1
2	Top down Einstein-Maxwell-pseudo-scalar model	5
3	Dyonic $AdS_2 \times \mathbb{R}^2$ solutions	6
4	Ansatz and thermodynamics	7
4.1	Ansatz for domain wall and black hole solutions	7
4.2	Asymptotic expansions	8
4.2.1	UV expansion	8
4.2.2	IR expansion: finite temperature	9
4.2.3	IR expansion: zero temperature	10
4.3	Thermodynamics	10
5	Dyonic domain wall solutions	12
6	Dyonic black hole solutions	13
6.1	Hyperscaling violation	16
6.2	Other dyonic black hole solutions	18
7	Superfluid and striped instabilities	18
7.1	Striped instabilities	18
7.2	Superfluid instabilities	19
7.3	Conclusions about the full phase diagram	21
8	Final comments	23
A	Dyonic AdS-RN black holes of Einstein-Maxwell theory	24

1 Introduction

The AdS/CFT correspondence allows us to study strongly coupled quantum field theories at finite temperature by studying appropriate black hole solutions of a dual gravitational theory. One focus has been the study of conformal field theories with AdS duals when held at finite chemical potential with respect to a global U(1) symmetry. Depending on the details of the gravitational theory, whose matter content includes a U(1) gauge-field, various types of novel phases are possible, corresponding to the existence of fascinating new classes of electrically charged black branes.

One well-studied possibility is that the CFT undergoes a superfluid phase transition. The first constructions were in the context of a bottom-up, Einstein-Maxwell theory of gravity coupled to a charged scalar field [1–3] and then extended to top-down constructions in [4, 5]. Another possibility is that the CFT undergoes a phase transition to a spatially modulated phase, in which translation invariance is spontaneously broken. This has been discussed in the context of electrically charged black holes [6–9] and also magnetically charged black holes [10, 11]. Other work, utilising the brane probe approximation, can be found in [12–15]. In some special examples in $D = 5$ it is possible to construct fully backreacted black hole solutions that are spatially modulated, with a helical structure, by solving ODEs [16, 17] (see also [18]). However, for most cases, including the $D = 4$ examples of interest in this paper, one will need to solve PDEs which is technically more challenging.¹

A simple diagnostic for the existence of new branches of black hole solutions at finite temperature can often be obtained by analysing the zero temperature limit of the unbroken, normal phase black hole solutions. Indeed, this is possible when the ground state that is approached in this limit has a finite entropy density, $s \neq 0$. For example, focussing on $D = 4$, if the zero temperature limit is described by a domain wall solution that interpolates between AdS_4 in the UV and $AdS_2 \times \mathbb{R}^2$ in the IR, as in the AdS-RN black brane solution, one can look for modes of the $AdS_2 \times \mathbb{R}^2$ solution that violate the AdS_2 BF bound. If such modes exist, the $T = 0$ domain wall solution will also be unstable and hence, by continuity, so will the finite-temperature black hole solutions, for low enough temperatures. To determine the critical temperature at which this instability sets in and where a new branch of black holes appears, one should look for static, normalisable, linearised perturbations about the finite temperature unbroken phase black hole solutions. While this approach is conceptually straightforward, one can encounter situations, as we will here, where the critical temperature is so low that it is hard to stabilise the numerical integration to find its precise value. One can nevertheless take the conceptual point of view that new phases emerge at low temperatures due to the destabilising effect of the finite ground-state entropy at $T = 0$ and thus that the qualitative features of the finite-temperature phase structure are dictated by the nature of the mode spectrum in the $s \neq 0$ ground state itself. One is thus naturally led to investigate the competition of possible sources of instabilities in the finite-entropy state.

In this paper we will employ these methods to analyse the competition between superfluid and spatially modulated or “striped”² phases [8] in a top-down setting. More specifically we will consider a $D = 4$ model in which both of these instabilities are present, and then switch on a magnetic field, aiming to suppress the superfluid instability while maintaining the striped instability. This scenario is in fact realised and we also find, for

¹The PDEs associated with some of the holographic striped black holes discovered in [8] were recently studied in [19].

²The word “striped” arises because the linearised mode for the spatially modulated branch has a striped structure, being translationally invariant along a given direction. At the linearised level one can also superpose such modes in different directions, losing this translation invariance. In order to determine the precise nature of the spatially modulated phase will require constructing the backreacted black holes.

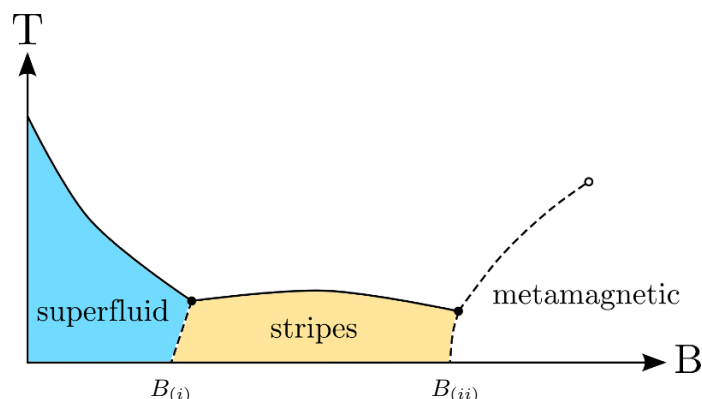


Figure 1. A schematic figure of a plausible phase diagram as a function of applied magnetic field. The solid lines denote second order phase transitions and the dashed lines first order. The two solid dots denote tri-critical points and the open circle a second order critical point where the metamagnetic transition ends. The phase diagram is symmetric under $B \rightarrow -B$. For $B > B_{(ii)}$ as $T \rightarrow 0$ the solutions exhibit hyperscaling violation in the IR with $z = 3/2$ and $\theta = -2$.

very large magnetic fields, a first-order metamagnetic phase transition at non-zero temperature. Such transitions, involving a discontinuous jump in the magnetisation of the system and not usually associated with symmetry breaking, are seen in a variety of materials, such as heavy fermion systems. They have also been discussed in a holographic context involving probe branes in [20], but as far as we are aware our construction with $T \neq 0$ is the first in a gravitational setting. Following these solutions down to zero temperature, we find that they exhibit hyperscaling violation in the far IR. Holographic hyperscaling violating solutions have been studied in [21–23] and top-down constructions appear³ in [25–28]. Our solutions become purely magnetic in the IR with dynamical exponent $z = 3/2$ and hyperscaling violation exponent $\theta = -2$. In particular, the entropy scales like $s \propto T^{8/3}$ as $T \rightarrow 0$.

A schematic picture of the likely phase diagram incorporating our findings is given in figure 1; there are some assumptions going into this figure which will be discussed in detail in the text. Notice that the metamagnetic transition ends in a critical point which is second-order (as in the liquid-vapour case) or higher, at finite temperature.

The $D = 4$ top-down model that we shall consider couples the metric with a gauge field, a charged scalar field and a neutral pseudo-scalar field, σ and arises from a consistent Kaluza-Klein (KK) reduction on an arbitrary $D = 7$ Sasaki-Einstein space SE_7 [24, 29]. This means that any solution of this $D = 4$ theory can be uplifted to $D = 11$ on an arbitrary SE_7 space leading to an infinite class of $D = 11$ solutions. In particular, the vacuum AdS_4 solution uplifts to the skew-whiffed $AdS_4 \times SE_7$ solutions which should⁴ be dual to $d = 3$ CFTs with an abelian global symmetry and, generically, no supersymmetry, apart from the special case when $SE_7 = S^7$ which preserves all of the supersymmetry.

³In fact the $T \rightarrow 0$ limit of the uncharged and also the charged normal phase black holes of [24] approach top-down hyperscaling violating solutions in the IR with $z = 1$ and $\theta = -1$.

⁴A dual CFT will exist provided that the skew-whiffed $AdS_4 \times SE_7$ solution is stable. It is known that skew-whiffed solutions are perturbatively stable [30]. A discussion of possible dual CFTs can be found in [31, 32].

The unbroken phase of these $d = 3$ CFTs at finite temperature, T , and chemical potential with respect to the abelian symmetry, μ , is described by the electrically charged AdS-RN black brane solution. A branch of superfluid black hole solutions appears at a critical temperature [33] and the fully back-reacted solutions were constructed in [24]. On the other hand at very low temperatures the AdS-RN black brane also connects with a spatially modulated “striped” branch of black holes [8]. This was demonstrated by constructing BF violating modes in the $AdS_2 \times \mathbb{R}^2$ solution arising in the IR limit of the $T = 0$ AdS-RN black hole solution. By analysing linearised modes about the finite temperature AdS-RN black holes, it was shown in [8] that the critical temperature for this transition is extremely low, but the precise value was not found. With these results we cannot be certain about what happens to the superfluid phase as it is cooled. The simplest possibility is that the system stays on the superfluid branch of black holes all the way down to zero temperature. However, it is possible that the system moves to a striped phase at low temperatures either via a first order or a second order phase transition. It is also possible that there are transitions to other phases. To establish this one would need to know, in principle, all of the black holes that exist at low temperatures, including the fully back reacted spatially modulated black holes. In this paper, we will not address these issues as they remain technically out of reach.

Here, instead, we will analyse the class of $d = 3$ skew-whiffed CFTs at finite T, μ after switching on a magnetic field, B . The presence of both electric and magnetic fields provides a source for the $D = 4$ neutral, pseudo-scalar field, σ , and this leads to a very rich structure for the unbroken, normal-phase black hole solutions (i.e. with vanishing charged scalar field). Indeed the study of these dyonic black hole solutions and related domain wall solutions that appear at zero temperature, both of which are solutions of a simple Einstein-Maxwell-pseudo-scalar theory in $D = 4$, and which are of interest in their own right, will be the focus for much of our analysis.⁵

We begin by showing that the $D = 4$ Einstein-Maxwell-pseudo-scalar model admits two families of dyonic $AdS_2 \times \mathbb{R}^2$ solutions carrying, generically, both electric and magnetic charges and $\sigma \neq 0$. One family, the “electric family”, contains the purely electrically charged $AdS_2 \times \mathbb{R}^2$ solution, arising in the electric AdS-RN solution at $T = 0$, while the other, the “magnetic family”, contains the purely magnetic solution, arising in the magnetic AdS-RN solution at $T = 0$. We then investigate which of these solutions can arise as the IR limit of a domain wall solution that asymptotes to AdS_4 in the UV, with deformation data given by μ, B . Such domain walls are possible zero temperature limits of finite temperature black holes describing the unbroken phase. Interestingly we find that there can be more than one domain wall solution with the same UV data μ, B and yet different $AdS_2 \times \mathbb{R}^2$ solutions in the IR. These solutions differ in the normalisable data in the UV and the parameters governing the irrelevant operators in the IR. Another surprising result is that while we find that some of the domain wall solutions can be heated up to arbitrarily high temperatures some cannot. Further insight into the unbroken phase can be obtained by

⁵It is worth comparing and contrasting our charged pseudo-scalar black holes with the charged, dilatonic black holes i.e. with neutral scalar fields, which have been studied in a holographic context in many works, including [21, 34–37].

calculating the free energy of the black holes and, in particular, we will see the first order metamagnetic phase transition appearing for large B . In contrast to the dyonic AdS-RN black holes of Einstein-Maxwell theory which are always strongly diamagnetic [3, 38] (as we review in appendix A), we will also see paramagnetic behaviour.

The plan of the rest of the paper is as follows. In section 2 we introduce the top-down Einstein-Maxwell pseudo-scalar model that will be used in sections 2–6. The new dyonic-pseudo-scalar $AdS_2 \times \mathbb{R}^2$ solutions are presented in section 3. In section 4 we introduce our ansatz for domain wall and black hole solutions and discuss some aspects of the thermodynamics. In particular we include a discussion of the definition of magnetisation and magnetic susceptibility. Sections 5 and 6 construct domain wall and black hole solutions, including some discussion of the magnetisation properties of the domain walls and the metamagnetic phase transition. Section 7 analyses the striped and superfluid instabilities of the $AdS_2 \times \mathbb{R}^2$ and the black hole solutions. In analysing the striped instabilities, which do not involve the charged scalar fields, we find the surprising result that, after a scaling of the wave-number, the spectrum of perturbations is the same for all dyonic $AdS_2 \times \mathbb{R}^2$ solutions. At the end of section 7 we summarise our conclusions about the full phase diagram, which lead to figure 1. We briefly conclude in section 8. Finally, we have one appendix where we calculate the magnetisation and susceptibility of dyonic black holes in Einstein-Maxwell theory to compare with the results that we obtain in our model.

2 Top down Einstein-Maxwell-pseudo-scalar model

For most of the paper we will consider the $D = 4$ model of [24] which couples the metric to a gauge-field, A , and a neutral pseudo-scalar field, σ . The action is given by

$$S = \frac{1}{16\pi G} \int d^4x \sqrt{-g} \left(R - \frac{1}{2}(\partial\sigma)^2 - \frac{\tau(\sigma)}{4}F^2 - V(\sigma) \right) + \frac{1}{32\pi G} \int \vartheta(\sigma)F \wedge F \quad (2.1)$$

where $F = dA$ and

$$V(\sigma) \equiv -24 \cosh \frac{\sigma}{\sqrt{3}}, \quad \tau(\sigma) \equiv \frac{1}{\cosh \sqrt{3}\sigma}, \quad \vartheta(\sigma) \equiv \tanh \sqrt{3}\sigma. \quad (2.2)$$

The inclusion of the charged field of [24] will be treated later when we discuss superfluid instabilities. Note that to compare⁶ with [24] we should set $h = \tanh \frac{\sigma}{\sqrt{3}}$ and also rescale the gauge field $A^{here} = 2A^{there}$. Occasionally we will find it convenient to use the field h instead of σ in some plots.

The equations of motion are given by

$$\begin{aligned} R_{ab} &= \frac{1}{2}\partial_a\sigma\partial_b\sigma + \frac{V}{2}g_{ab} + \frac{\tau}{2}\left(F_{ab}^2 - \frac{1}{4}g_{ab}F^2\right), \\ d(\tau * F) &= d\vartheta \wedge F, \\ d * d\sigma &= -V' * 1 - \frac{\tau'}{2}F \wedge *F + \frac{\vartheta'}{2}F \wedge F. \end{aligned} \quad (2.3)$$

⁶To compare with [8] we should set $\sigma = \sqrt{2}\varphi$, rescale the metric $g^{here} = (1/2)g^{there}$, identify the potentials $V^{here} = 4V^{there}$ and also $\vartheta^{here} = -\vartheta^{there}$.

Any solution of these equations of motion, for the specific functions given in (2.2), can be uplifted on an arbitrary SE_7 manifold to obtain an exact solution of $D = 11$ supergravity using the formulae in [24]. For example the basic vacuum AdS_4 solution with $\sigma = A = 0$ and radius squared $1/4$ uplifts to the standard skew-whiffed $AdS_4 \times SE_7$ solution, which doesn't preserve any supersymmetry except in the special case of $SE_7 = S^7$ in which it preserves all of the supersymmetry.

It is interesting to observe that given⁷ (2.2), the equations of motion exhibit the following electric/magnetic duality symmetry

$$F \rightarrow -\tau(\sigma) * F + \vartheta(\sigma)F, \quad \sigma \rightarrow -\sigma \tag{2.4}$$

The origin of this symmetry should be related to the fact that the action comes from a truncation of an $N = 2$ supergravity theory [24, 29]. We also observe that, for the specific functions given in (2.2), it is only consistent to set $\sigma = 0$ for configurations that have $F \wedge F = 0$. For such configurations, the equations of motion collapse to those of Einstein-Maxwell theory. We will concentrate on the functions given in (2.2) but many of our results have simple generalisations for different choices of functions.

3 Dyonic $AdS_2 \times \mathbb{R}^2$ solutions

We consider the following ansatz

$$\begin{aligned} ds^2 &= L^2 ds^2(AdS_2) + ds^2(\mathbb{R}^2), \\ F &= -EL^2 \text{Vol}(AdS_2) + B \text{Vol}(\mathbb{R}^2), \end{aligned} \tag{3.1}$$

where $ds^2(AdS_2)$ and $\text{Vol}(AdS_2)$ are the metric and volume form for a unit radius AdS_2 , and σ, L, E and B are constants, and the minus sign appearing is for later convenience. Substituting into the equations of motion for (4.2) we are led to the following algebraic conditions

$$E^2 + B^2 = -\frac{2V}{\tau}, \quad \frac{\tau'}{2}(E^2 - B^2) - \vartheta'EB - V' = 0, \tag{3.2}$$

with the AdS_2 radius given by

$$L^{-2} = -V. \tag{3.3}$$

Notice that these equations are invariant under simultaneously flipping the sign of B and σ . In addition the duality transformation (2.4) corresponds to

$$E \rightarrow \tau B + \vartheta E, \quad B \rightarrow -\tau E + \vartheta B, \quad \sigma \rightarrow -\sigma \tag{3.4}$$

Up to flipping the sign of both E and B , for the specific functions given in (2.2), these relations define two one-parameter families of dyonic $AdS_2 \times \mathbb{R}^2$ solutions, labelled by σ ,

⁷More generally, we require that (τ, V) and ϑ are even and odd functions of σ , respectively, that are analytic at $\sigma = 0$.

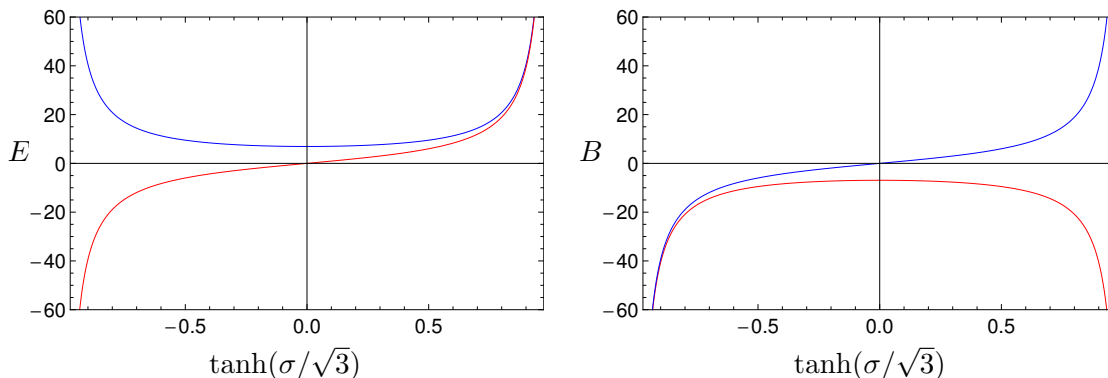


Figure 2. Two families of dyonic $AdS_2 \times \mathbb{R}^2$ solutions. The left and right panels display the dependence of E and B on the value of the pseudo-scalar σ , respectively. For convenience of presentation we use $\tanh(\sigma/\sqrt{3})$ for the horizontal axis. The blue line is the electric family which contains a purely electric solution and the red line is the magnetic family which contains a purely magnetic solution. There are two more families of solutions obtained by simultaneously flipping the signs of E, B .

which we have summarised in figure 2. Notice that one family, which we call the “electric family”, contains a purely electric solution, while the other “magnetic family”, contains a purely magnetic solution. As we discuss below, the purely electric and magnetic $AdS_2 \times \mathbb{R}^2$ solutions, both of which have $\sigma = 0$, arise as the near horizon limit of the standard electric and magnetic AdS-RN black brane solutions, respectively.

4 Ansatz and thermodynamics

4.1 Ansatz for domain wall and black hole solutions

In the sequel we will consider the following ansatz

$$\begin{aligned}
 ds^2 &= -e^{-\beta(r)}g(r)dt^2 + \frac{dr^2}{g(r)} + r^2(dx^2 + dy^2), \\
 A &= \phi(r)dt + B\frac{1}{2}(xdy - ydx), \\
 \sigma &= \sigma(r).
 \end{aligned}
 \tag{4.1}$$

After substituting into the equations of motion (2.3) we obtain differential equations for the functions β, g, ϕ and σ . These equations can also be obtained by substituting the ansatz directly into the action (2.1), leading to

$$\begin{aligned}
 S = c_0 \int dr r^2 e^{-\beta/2} &\left[-g'' + g' \left(\frac{3}{2}\beta' - \frac{4}{r} \right) + g \left(\beta'' - \frac{1}{2}(\beta')^2 + 2\frac{\beta'}{r} - \frac{2}{r^2} \right) \right. \\
 &\left. - \frac{1}{2}g\sigma'^2 + \frac{1}{2}\tau(\sigma) \left(e^\beta \phi'^2 - \frac{B^2}{r^4} \right) - V(\sigma) - B\vartheta(\sigma) \frac{e^{\beta/2}\phi'}{r^2} \right],
 \end{aligned}
 \tag{4.2}$$

where $c_0 = (16\pi G)^{-1} \int dt dx dy$, and then varying. It is helpful to observe that the ansatz and hence the equations of motion are invariant under the two scalings

$$\begin{aligned} t &\rightarrow \lambda t, & e^\beta &\rightarrow \lambda^2 e^\beta, & \phi &\rightarrow \lambda^{-1} \phi; \\ r &\rightarrow \lambda r, & (t, x, y) &\rightarrow \lambda^{-1}(t, x, y), & g &\rightarrow \lambda^2 g, & \phi &\rightarrow \lambda \phi, & B &\rightarrow \lambda^2 B. \end{aligned} \quad (4.3)$$

We also note the symmetry

$$B \rightarrow -B, \quad \sigma \rightarrow -\sigma \quad (4.4)$$

When $\sigma = 0$, the standard electrically charged AdS-RN black brane solves the equations of motion. It is given by

$$g = 4r^2 - \left(4r_+^2 + \frac{\mu^2}{4}\right) \frac{r_+}{r} + \frac{\mu^2 r_+^2}{4r^2}, \quad \phi = \mu \left(1 - \frac{r_+}{r}\right) \quad (4.5)$$

with $\beta = B = 0$. When $\mu = 4\sqrt{3}r_+$, the temperature $T = 0$ and the solution is a domain wall interpolating between AdS_4 in the UV and the purely electrically charged $AdS_2 \times \mathbb{R}^2$ solution in the IR given in (3.1), (3.2), (3.3) with $L^2 = 1/24$ and $E = -4\sqrt{3}$. Although we will be principally interested in black hole solutions with non-vanishing electric charge, we recall here the purely magnetic AdS-RN black brane solution:

$$g = 4r^2 - \left(4r_+^2 + \frac{B^2}{4r_+^2}\right) \frac{r_+}{r} + \frac{B^2}{4r^2}, \quad (4.6)$$

with $\beta = \phi = 0$. The $T = 0$ limit is when $B = 4\sqrt{3}r_+^2$ and the solution becomes a domain wall approaching the purely magnetically charged $AdS_2 \times \mathbb{R}^2$ solution in the IR given in (3.1), (3.2), (3.3) with $L^2 = 1/24$ and $B = 4\sqrt{3}$.

Notice that the standard dyonic AdS-RN black hole is not a solution to the equations of motion, since the presence of both electric and magnetic fields sources the pseudo-scalar σ . We will construct dyonic black hole and domain wall solutions with $\sigma \neq 0$ numerically in later sections.

4.2 Asymptotic expansions

4.2.1 UV expansion

We are interested in studying the $d = 3$ CFTs with chemical potential μ and magnetic field B . In the UV, as $r \rightarrow \infty$, we will impose the following expansions

$$\begin{aligned} g &= 4r^2 + \sigma_1^2 - \frac{1}{2r} (\varepsilon - 4\sigma_1\sigma_2) + \dots, \\ \beta &= \beta_a + \frac{\sigma_1^2}{4r^2} + \frac{2\sigma_1\sigma_2}{3r^3} \dots, \\ \phi &= e^{-\frac{\beta_a}{2}} \left(\mu - \frac{q}{r} - \frac{\sqrt{3}B\sigma_1}{2r^2} + \dots \right), \\ \sigma &= \frac{\sigma_1}{r} + \frac{\sigma_2}{r^2} + \frac{5\sigma_1^3}{72r^3} \dots \end{aligned} \quad (4.7)$$

Note that for simplicity⁸ we focus on the quantisation of the pseudo-scalar so that the dual operator has dimension $\Delta(\mathcal{O}_\sigma) = 2$. For the most part, we will consider the CFT with no deformation in the UV by \mathcal{O}_σ corresponding to setting $\sigma_1 = 0$. For simplicity of presentation, we have set

$$16\pi G = 1 \tag{4.8}$$

The appropriate factors can easily be reinstated if required. It is also worth repeating here that the radius of the asymptotic AdS_4 is $1/2$.

4.2.2 IR expansion: finite temperature

In this case we are interested in black hole solutions with regular event horizons located at $r = r_+$ where there is an analytic expansion of the form

$$\begin{aligned} g &= g_+(r - r_+) + \dots, \\ \beta &= \beta_+ + \dots, \\ \phi &= \phi_+(r - r_+) + \dots, \\ \sigma &= \sigma_+ + \dots \end{aligned} \tag{4.9}$$

This expansion is fixed by 4 constants, $\beta_+, \phi_+, \sigma_+$ and r_+ with, for example,

$$g_+ = 12r_+ \cosh \frac{\sigma_+}{\sqrt{3}} - \frac{B^2 + r_+ e^{\beta_+} \phi_+^2}{4 \cosh \sqrt{3} \sigma_+}. \tag{4.10}$$

A black hole solution is then specified by 6 UV parameters (with $\sigma_1 = 0$) and 4 IR parameters. From (4.2) we have two first order equations of motion, for g, β , and two second order equations, for ϕ, σ and so a solution is specified by 6 integration constants. Taking into account the scaling symmetries (4.3), we expect two-parameter families of black hole solutions which can be labelled by temperature T and magnetic field B , or better, by the dimensionless quantities T/μ and B/μ^2 .

For later use we notice that the equation of motion for ϕ arising from (4.2) (i.e. the equation of motion for the gauge-field in (2.3)) can be integrated from $r = r_+$ to $r = \infty$ and from (4.9) and (4.7) we obtain the charge conservation condition

$$q = \tau(\sigma_+) e^{\beta_+/2} \phi_+ r_+^2 - B \vartheta(\sigma_+), \tag{4.11}$$

with the first and second terms on the right hand side arising from the F^2 term and the $F \wedge F$ terms in the action (2.1), respectively. This condition can be satisfied in the far IR of a black hole solution in the zero temperature limit in different ways. For example, we will construct dyonic black hole solutions which approach domain walls at $T = 0$ with non-zero entropy density, which will get contributions from both terms. We will also construct black hole solutions with vanishing entropy density as $T \rightarrow 0$ with all of the contribution to q coming from the $F \wedge F$ term.

⁸For the special case of uplifting the $D = 4$ solutions on an S^7 , the choice $\Delta(\mathcal{O}_\sigma) = 2$ is associated with maximal supersymmetry.

4.2.3 IR expansion: zero temperature

We are also interested in domain wall solutions that asymptote in the IR, as $r \rightarrow r_+$, to the dyonic $AdS_2 \times \mathbb{R}^2$ solutions given in (3.1)–(3.3). We focus on the following expansion

$$\begin{aligned} g(r) &= g_0 (r - r_+)^2 + \dots \\ \sigma(r) &= \sigma_0 + \dots + \sigma_+ (r - r_+)^{\Delta-1} + \dots \\ \phi(r) &= \phi_0 (r - r_+) + \dots + \sigma_+ \delta\phi (r - r_+)^{\Delta} + \dots \\ \beta(r) &= \beta_0 + \dots + \sigma_+ \delta\beta (r - r_+)^{\Delta-1} + \dots \end{aligned} \tag{4.12}$$

where g_0 is determined via

$$g_0 = -V(\sigma_0), \tag{4.13}$$

together with the relations

$$E^2 + \tilde{B}^2 = -\frac{2V(\sigma_0)}{\tau(\sigma_0)}, \quad \frac{\tau'(\sigma_0)}{2}(E^2 - \tilde{B}^2) - \vartheta'(\sigma_0)E\tilde{B} - V'(\sigma_0) = 0, \tag{4.14}$$

where $E \equiv e^{\frac{\beta_0}{2}} \phi_0$ and $\tilde{B} \equiv B/r_+^2$. Also, σ_+ parametrises a deformation by an irrelevant operator of dimension Δ , with

$$\Delta = \frac{1}{6}(3 + \sqrt{105}), \tag{4.15}$$

which we discuss further below. As $r \rightarrow r_+$ this expansion approaches the exact dyonic $AdS_2 \times \mathbb{R}^2$ solutions given in (3.1)–(3.3). The irrelevant deformation with Δ as (4.15) is obtained by linearising about an exact dyonic $AdS_2 \times \mathbb{R}^2$ solution. We find that the corresponding mode has

$$\begin{aligned} \delta\beta &= -2\sqrt{3} \tanh \frac{\sigma_0}{\sqrt{3}}, \\ \delta\phi &= e^{-\frac{\beta_0}{2}} \frac{\sqrt{3} \tilde{B} + 2E \sinh \sqrt{3}\sigma_0 - 2E \sinh \frac{\sigma_0}{\sqrt{3}}}{2\Delta \cosh \sqrt{3}\sigma_0}. \end{aligned} \tag{4.16}$$

A domain wall solution is specified by 6 UV parameters (with $\sigma_1 = 0$) and 3 IR parameters (i.e. $\sigma_0, \phi_0, \beta_0, r_+, \sigma_+$ subject to the two constraints (4.14)). Now from (4.2) we have two first order equations of motion, for g, β , and two second order equations, for ϕ, σ and so a solution is specified by 6 integration constants. Taking into account the scaling symmetries (4.3), we expect a one-parameter family of black hole solutions which can be labelled by the magnetic field B .

4.3 Thermodynamics

Building on [3, 39], we generalise the discussion of [24] to include $B \neq 0$. We analytically continue by setting $t = -i\tau$, together with $I = -iS$. We can then obtain two expressions for the on-shell action for the class of solutions we are studying. The first expression is given by the integral of a total derivative

$$I_{OS} = \Delta \tau vol_2 \int dr \left[r^2 e^{-\frac{\beta}{2}} \left(g' - g\beta' - \tau(\sigma) e^{\beta} \phi \phi' \right) + B \vartheta(\sigma) \phi \right]'. \tag{4.17}$$

The second can be written

$$I_{OS} = \Delta\tau vol_2 \int dr \left\{ \left[2rge^{-\beta/2} \right]' + B^2 r^{-2} e^{-\beta/2} \tau(\sigma) + B\vartheta(\sigma)\phi' \right\}, \quad (4.18)$$

which is an integral of a total derivative only in the special case that $B = 0$.

We define the total action I_{Tot} via

$$I_{Tot} = I + I_{ct}, \quad (4.19)$$

where the boundary counter term action is

$$I_{ct} = \int d\tau d^2x \sqrt{g_\infty} (-2K + 8 + \sigma^2). \quad (4.20)$$

We next define the thermodynamic potential $W \equiv T[I_{tot}]_{OS} \equiv w vol_2$ where the temperature of the black hole is given by

$$T = \frac{e^{\beta_a/2}}{4\pi} [g' e^{-\beta/2}]_{r=r_+}. \quad (4.21)$$

Corresponding to the expression (4.17), and using the expansion (4.7) as well as (4.9), we obtain

$$w = \varepsilon - \mu q - T s. \quad (4.22)$$

On the other hand, corresponding to (4.18) we find the alternative expression

$$w = -\left(\frac{\varepsilon}{2} + 2\sigma_1\sigma_2\right) + e^{\beta_a/2} \int dr \left\{ B^2 r^{-2} e^{-\beta/2} \tau(\sigma) + B\vartheta(\sigma)\phi' \right\}. \quad (4.23)$$

The equality of these two expressions gives a Smarr type formula, which we shall return to below.

We next consider an on-shell variation of the total action, as in [24], and deduce that $w = w(T, \mu, \sigma_1, B)$ with

$$\delta w = -s\delta T - q\delta\mu - 4\sigma_2\delta\sigma_1 - m\delta B, \quad (4.24)$$

where the entropy density, s , is given by

$$s = 4\pi r_+^2, \quad (4.25)$$

and the magnetisation per unit volume, $m \equiv -\partial w / \partial B$ at constant T, μ, σ_1 , is given by

$$m = -e^{\beta_a/2} \int dr \left\{ B r^{-2} e^{-\beta/2} \tau(\sigma) + \vartheta(\sigma)\phi' \right\}. \quad (4.26)$$

Observe that the axionic-like coupling ϑ in (2.1) can give rise to magnetisation even when $B = 0$.

The holographic stress tensor of [40], given by

$$T_i^j = (2r^3)[-2K_i^j + 2\delta_i^j(2K - 8 - \sigma^2)], \quad (4.27)$$

can be calculated and we find

$$\begin{aligned}
 T_t^t &= -\varepsilon, \\
 T_x^x = T_y^y &= \frac{\varepsilon}{2} + 2\sigma_1\sigma_2.
 \end{aligned}
 \tag{4.28}$$

Thus ε is the energy of the system. We also see that when $\sigma_1 = 0$, corresponding to no deformation by \mathcal{O}_σ , but with $B, \mu \neq 0$, the stress tensor is traceless. Using the general result that $w = -p$, where p is the pressure, a comparison of (4.23), (4.26) (4.28) reveals that the content of the Smarr relationship can be written

$$\begin{aligned}
 p &= \frac{\varepsilon}{2} + 2\sigma_1\sigma_2 + mB, \\
 &= T_x^x + mB,
 \end{aligned}
 \tag{4.29}$$

similar to what was noted in [3]. The magnetic susceptibility, χ_m , is defined by $\chi_m = \partial m / \partial B = -\partial^2 w / \partial B^2$ with the derivatives taken at constant T, μ, σ_1 .

5 Dyonic domain wall solutions

In this section we construct domain wall solutions that approach AdS_4 in the UV as in (4.7), and we set $\sigma_1 = 0$ corresponding to no deformation of the CFT by the operator \mathcal{O}_σ . We also use the scaling (4.3) to set $\beta_a = 0$. In the IR they approach one of the dyonic $AdS_2 \times \mathbb{R}^2$ solutions discussed in section 3 via the expansion (4.12). As discussed above, we expect a one-parameter family of solutions which we label by the value of the dimensionless quantity B/μ^2 . For definiteness we will choose $\mu > 0$ and focus on $B \geq 0$. For $\sigma_1 = 0$, the solutions with $B \leq 0$ can be obtained using the symmetry (4.4).

We first consider domain wall solutions that in the IR approach $AdS_2 \times \mathbb{R}^2$ solutions lying in the electric family. The simplest case is the purely electric $AdS_2 \times \mathbb{R}^2$ solution for which the unique domain wall solution is simply the $T = 0$ limit of the electrically charged AdS-RN black brane (4.5). As we switch on B we find a one-parameter family of solutions as given in figure 3. The first interesting feature is that the domain wall solutions only exist up to a maximum value of B/μ^2 given by $(B/\mu^2)_{\max} \approx 0.12$. The second interesting feature is that for fixed values of $B/\mu^2 \neq 0$ there can be two distinct domain wall solutions. These domain walls have the same UV deformation data μ, B but are distinguished in having different values of ε, q, σ_2 in the UV expansion (4.7), corresponding to different values of the energy, charge and $\langle \mathcal{O}_\sigma \rangle$. After calculating the free energy w for these solutions we find that the upper branch in the left panel in figure 3 is always thermodynamically preferred. The numerical results suggest that this picture persists for all values of $B/\mu^2 < (B/\mu^2)_{\max}$. Figure 3 also displays the magnetisation m for these domain wall solutions. It is interesting to observe that it is always positive, corresponding to a paramagnetic system, in contrast to the dyonic AdS-RN solutions of pure Einstein-Maxwell theory which are diamagnetic [3, 38] (see appendix A).

We next consider domain wall solutions that approach $AdS_2 \times \mathbb{R}^2$ solutions lying in the magnetic family. We display a one-parameter family in figure 4 again with $B > 0, \mu > 0$.

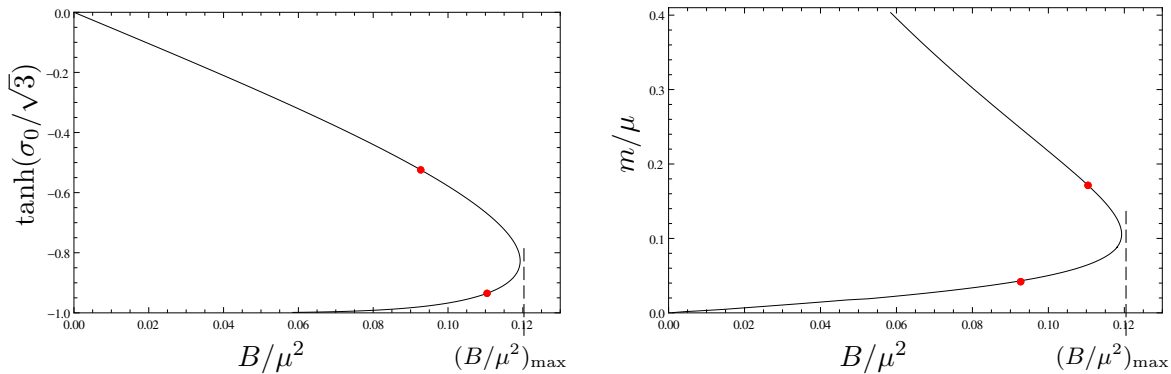


Figure 3. Left panel: the one parameter family of domain wall solutions interpolating between AdS_4 in the UV, with deformation data (μ, B) , and dyonic $AdS_2 \times \mathbb{R}^2$ solutions in the electric family in the IR, labelled by σ_0 . For convenience of presentation the vertical axis is given by $\tanh(\sigma_0/\sqrt{3})$. There can be two domain wall solutions for given (μ, B) and the upper branch has smaller free energy and is thermodynamically preferred. It is expected that the lower branch continues down to $B/\mu^2 \rightarrow 0$. The red dots indicate superfluid instabilities, discussed in section 7.2, with the solutions being unstable to the left of the dots. Right panel: a plot of the magnetisation m/μ as a function of B/μ^2 . Observe that the magnetisation is always positive corresponding to paramagnetism.

Note that the purely magnetic solution, the $T = 0$ limit of the purely magnetically charged AdS-RN solution (4.6), is obtained when $\mu = 0$ and hence $B/\mu^2 \rightarrow \infty$ in the figure. Also note that we have only been able to stabilise the numerics for values of $B/\mu^2 \gtrsim 1$ and in particular we have not obtained these solutions for values of B/μ^2 that overlap with the solutions in figure 3. It is most likely that they do exist for all values of B/μ^2 down to zero, but that they have higher free energy and hence are not thermodynamically relevant.⁹ We will continue with this assumption. Figure 4 also displays the magnetisation m for these domain wall solutions. It is interesting to observe that it is always negative, corresponding to diamagnetism, in contrast to the domain wall solutions lying in the electric family of figure 3.

6 Dyonic black hole solutions

In this section we construct finite temperature black hole solutions that approach AdS_4 in the UV as in (4.7), again with $\sigma_1 = \beta_a = 0$. Also, as in the previous section, we choose $\mu > 0$ and $B \geq 0$ with solutions with $B \leq 0$ obtained using the symmetry (4.4).

Our initial strategy is to heat up the $T = 0$ domain wall solutions that we constructed in the previous section. One focus of interest is what we will call “region I”, with $0 \leq B/\mu^2 \leq (B/\mu^2)_I$. In this region, by definition, the unbroken phase is described by dyonic black holes whose zero temperature limit is given by dyonic domain wall solutions approaching

⁹This is based on two calculations. Firstly, using the values for the free energy of the domain walls that we have constructed in figure 4 and then extrapolating to smaller values of B/μ^2 . And secondly, we have also calculated some domain walls after adding in a deformation in the UV by the $\Delta = 2$ operator dual to σ (see section 8) where we were able to make a direct comparison of the free energies for the same UV data in some cases.

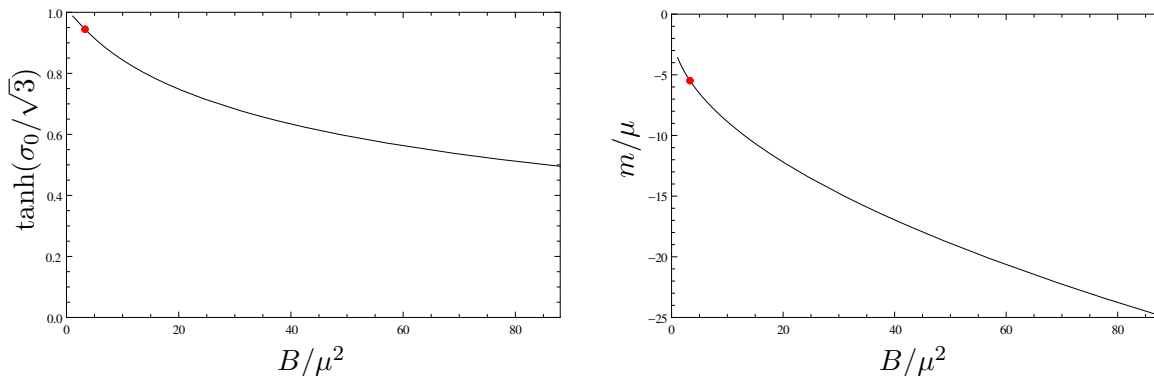


Figure 4. Left panel: the one parameter family of domain wall solutions interpolating between AdS_4 in the UV, with deformation data (μ, B) , and dyonic $AdS_2 \times \mathbb{R}^2$ solutions in the magnetic family in the IR, labelled by σ_0 . For convenience of presentation the vertical axis is given by $\tanh(\sigma_0/\sqrt{3})$. It is expected that these domain walls exist for $B/\mu^2 \rightarrow 0$ but that they have higher free energy than the domain walls in figure 3 for the same values of B, μ . The red dots indicate superfluid instabilities, discussed in section 7.2, with the solutions being unstable to the left of the dots. Right panel: a plot of the magnetisation m/μ as a function of B/μ^2 . Observe that the magnetisation is always negative corresponding to diamagnetism.

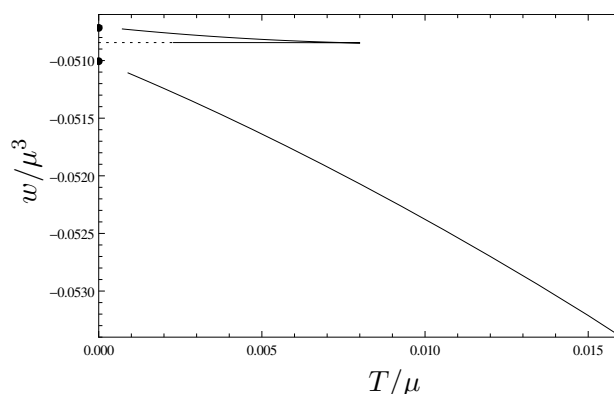


Figure 5. A representative plot of the free energy of two families of dyonic black hole solutions that exist in region I i.e. for values of B/μ^2 with $0 < B/\mu^2 < (B/\mu^2)_I$. At $T = 0$ each family approaches a smooth domain wall solution interpolating between AdS_4 in the UV and two different dyonic $AdS_2 \times \mathbb{R}^2$ solutions in the IR. Notice that only the bottom branch can be heated up to arbitrarily high temperatures; it is thermodynamically preferred and describes the unbroken phase.

in the IR $AdS_2 \times \mathbb{R}^2$ solutions lying on the electric family. One might suspect that $(B/\mu^2)_I$ coincides with $(B/\mu^2)_{\max}$, the value of B/μ^2 in which such domain walls cease to exist (see figure 3). However, this is not quite true and in fact we have $(B/\mu^2)_I < (B/\mu^2)_{\max}$, because of the existence of a first order metamagnetic transition, as we shall explain.

In figure 5 we show a representative plot of the black hole solutions that we have constructed with $0 < B/\mu^2 < (B/\mu^2)_I$, with B/μ^2 close to $(B/\mu^2)_I$. We can heat up the domain wall solution approaching the $AdS_2 \times \mathbb{R}^2$ solution that lies on the upper branch of figure 3 to arbitrary high temperatures. On the other hand for the domain wall ap-

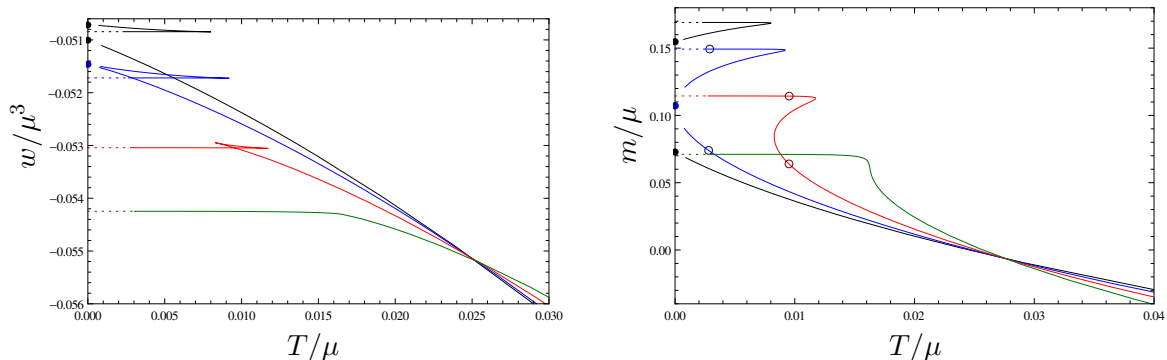


Figure 6. A series of dyonic black holes for different values of B/μ^2 , with increasing values of B/μ^2 for the black, blue, red and green curves, respectively. The black curve has $B/\mu^2 < (B/\mu^2)_I$. The blue curve has $B/\mu^2 = (B/\mu^2)_{\max} > (B/\mu^2)_I$ and the red and green curves have $B/\mu^2 > (B/\mu^2)_{\max}$. In the left and right plots we show the free energy and the magnetisation as a function of T . Observe that for the blue and red curves there is a first order phase transition at finite temperature, marked with circles on the right plot, where there is a discontinuous jump in the magnetisation. The solid black and blue dots refer to domain wall solutions with dyonic $AdS_2 \times \mathbb{R}^2$ solutions in the IR. The dashed curves indicate black hole solutions that are approaching hyperscaling violating behaviour in the IR as $T \rightarrow 0$.

proaching the $AdS_2 \times \mathbb{R}^2$ solution that lies on the lower branch of figure 3 we find that it can only be heated up to a maximum temperature before it goes back down to lower temperatures. Following this solution down to very low temperatures we find that the solution is becoming singular in the IR, approaching a hyperscaling violating behaviour as we discuss below. We can see from figure 5 that the free energy for the branch of black holes that can be heated up to high temperatures is always thermodynamically preferred and hence describes the high temperature unbroken phase of the system. In the right plot of figure 6 the black curves plot the magnetisation as a function of T for the two black hole branches for the value of B/μ^2 given in figure 5 (the same value of B/μ^2 for the black curve in the left plot of figure 6). We observe that for the thermodynamically preferred black holes the behaviour switches from paramagnetism at low temperatures to diamagnetism at high temperatures.

Having discussed the black holes for a representative value of $B/\mu^2 < (B/\mu^2)_I$ in region I, let us see what happens as we increase the value of B/μ^2 . A series of black hole solutions is presented in figure 6. We see that at $B = B_{\max}$, corresponding to the blue curve, the two branches of black holes are coalescing at $T = 0$ corresponding to the fact that at $B/\mu^2 = (B/\mu^2)_{\max}$ there is just a single domain wall solution mapping in the IR onto an $AdS_2 \times \mathbb{R}^2$ solution in the electric family. By examining the free energy, we see that the blue curve implies there is a first order phase transition at finite temperature. We also see from the right panel in figure 6, that there is an abrupt change in the magnetisation and hence we have a first-order metamagnetic phase transition. At low temperatures the preferred black holes, denoted by the dashed blue line in figure 6, all exhibit hyperscaling violation in the IR, as we elaborate on further below.

For higher values of B/μ^2 , $B/\mu^2 > (B/\mu^2)_{\max}$, illustrated by the red and the green curves in figure 6, there is just a single branch of black holes all of which approach a singular

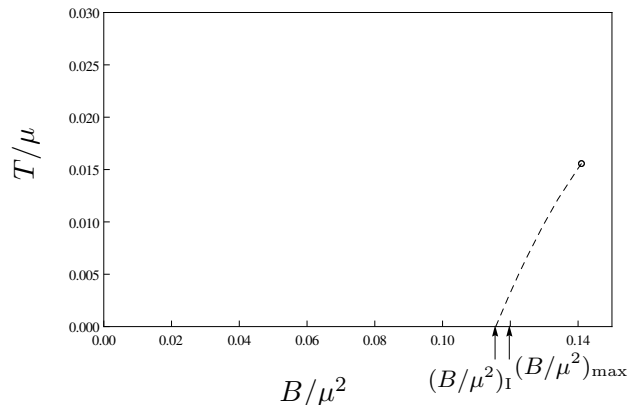


Figure 7. The phase diagram for the unbroken phase black holes. For Region I, with $0 \leq B/\mu^2 < (B/\mu^2)_I$ the black holes at $T = 0$ are domain walls that approach an $AdS_2 \times \mathbb{R}^2$ solution in the electric family in the far IR. More precisely, they are domain walls that are on the upper part of the curve in the left hand plot in figure 3. For $(B/\mu^2)_I \leq B/\mu^2 \leq (B/\mu^2)_{\max}$ similar black holes exist, but they are not thermodynamically preferred. Instead the dashed line represents a line of first order phase transitions to a metamagnetic phase, ending in a second order critical point denoted by a round circle. For $B/\mu^2 > (B/\mu^2)_I$, as $T \rightarrow 0$ the thermodynamically preferred solutions exhibit hyperscaling violation in the IR with dynamical exponent $z = 3/2$ and $\theta = -2$.

solution at $T = 0$. For the red curve we again see that there is a first order phase transition at finite temperature with the transition temperature depending on B/μ^2 . Increasing B/μ^2 further we get to the green curve and the first order transition comes to an end at a critical point which is second (or higher) order. In the right plot of figure 6 we see that the jump in the magnetisation, present for the first order transitions, is decreasing as one increases B/μ^2 , disappearing for the green curve. In figure 7 we have summarised the T, B phase diagram for these unbroken phase black holes.

6.1 Hyperscaling violation

We now return to the properties of the solutions denoted by dashed curves in figure 6 in the limit that $T \rightarrow 0$. We first observe from the free-energy curves in figure 6 that as $T \rightarrow 0$ they all have $\partial_T w \rightarrow 0$ and hence the entropy density $s \rightarrow 0$. In figure 8 we have plotted the behaviour of $\log s$ versus $\log T$ which clearly reveals, especially for large values of B/μ^2 , that $s \propto T^{8/3}$ as $T \rightarrow 0$. This strongly suggests that the solutions all exhibit an emergent scaling behaviour in the IR as $T \rightarrow 0$.

The scaling behaviour can be identified by returning to the equations of motion coming from (4.2). In particular, we have found the following one-parameter family of solutions, at leading order in r as $r \rightarrow 0$, given by

$$\begin{aligned}
 ds^2 &= -r^{5/2} dt^2 + \frac{dr^2}{2Lr} + r^2(dx_1^2 + dx_2^2), \\
 F &= B dx_1 \wedge dx_2, \\
 \sigma &= \sigma_0 + \sqrt{3} \log r,
 \end{aligned}
 \tag{6.1}$$

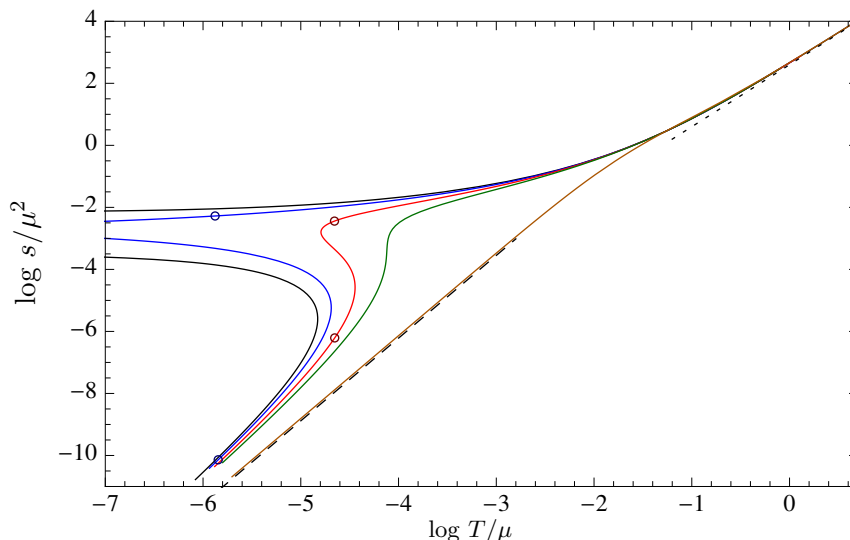


Figure 8. A plot of $\log s/\mu^2$ versus $\log T/\mu$ for the dyonic black hole solutions with different values of B/μ^2 . The black, blue, red and green curves are the same black holes as in figure 6 and the brown curve is a black hole with a much larger value of B/μ^2 . As $T \rightarrow \infty$ the dotted line corresponds to $s \propto T^2$, associated with the AdS_4 asymptotics in the UV. As $T \rightarrow 0$ the dashed line corresponds to $s \propto T^{8/3}$, associated with the hyperscaling violating asymptotics in the IR with $z = 3/2$ and $\theta = -2$.

where σ_0 is a constant and

$$B = \pm \frac{2}{\sqrt{3}} e^{-\frac{2\sigma_0}{\sqrt{3}}}, \quad L = \frac{32}{33} e^{-\frac{\sigma_0}{\sqrt{3}}}. \quad (6.2)$$

This purely magnetic configuration solves the equations of motion to leading order as $r \rightarrow 0$, and moreover, we find that the $F \wedge F$ coupling in (2.1) is not playing a role. A simple co-ordinate transformation $r \rightarrow (L/2)\rho^{-2}$, combined with a rescaling of the time and spatial co-ordinates, reveals this to be a hyperscaling violating metric [21–23] given by

$$ds^2 = \rho^{-(2-\theta)} (-\rho^{-2(z-1)} dt^2 + d\rho^2 + d\bar{x}_i dx_i), \quad (6.3)$$

with dynamical exponent $z = 3/2$ and hyperscaling violation exponent $\theta = -2$. In particular, under the scaling

$$t \rightarrow \lambda^z t, \quad x_i \rightarrow \lambda x_i, \quad \rho \rightarrow \lambda \rho, \quad (6.4)$$

the metric scales as $ds \rightarrow \lambda^{\theta/2} ds$. If one heats up this class of hyperscaling violating solutions one finds that the entropy density behaves as $s \propto T^{(2-\theta)/z} = T^{8/3}$, which is exactly the same behaviour we see for our solutions as $T \rightarrow 0$.

A more detailed look at our numerical solutions provides additional evidence that as $T \rightarrow 0$ the solutions are domain walls interpolating between the hyperscaling violating behaviour (6.1) in the IR and AdS_4 in the UV. We leave the detailed construction of such domain wall solutions to future work, but we note one final point. As $T \rightarrow 0$ the black

hole solutions have $\sigma_+ \rightarrow -\infty$ at the event horizon and also the UV charge q is given by $q \rightarrow B$. A consideration of (4.11) shows that the origin of the electric charge in the far IR for these $s = 0$ ground states is arising purely from the $F \wedge F$ coupling in (2.1).

6.2 Other dyonic black hole solutions

The alert reader will have noticed that in addition to the two branches of black holes presented in figure 5, mapping onto a $T = 0$ domain wall solution approaching an $AdS_2 \times \mathbb{R}^2$ solution in the electric family, there could be an additional branch of black holes which maps onto a $T = 0$ domain wall solution approaching an $AdS_2 \times \mathbb{R}^2$ solution in the magnetic family as in figure 4. Indeed we have constructed such solutions for values of B/μ^2 as in figure 4, which we recall are much higher than those in figure 5. By extrapolating the free energy of these solutions down to smaller values of B/μ^2 , combined with the observations that we made in footnote 6, we expect these black holes are never thermodynamically preferred over those presented in figure 5 and hence they will not change the picture summarised in figure 7. We also note that the solutions of this type that we have constructed exist up to a maximum temperature and then return to low temperatures, approaching a singular solution as $T \rightarrow 0$, analogous to the upper branch in figure 5. In fact these solutions also appear to approach the hyperscaling violating solutions given in (6.1).

7 Superfluid and striped instabilities

In section 7.1 we analyse striped instabilities of the dyonic $AdS_2 \times \mathbb{R}^2$ solutions in both the electric and magnetic families given in section 3. This allows us to deduce the existence of instabilities of the corresponding domain walls constructed in section 5 and the thermodynamically preferred unbroken phase black hole solutions in region I, i.e. with $0 \leq B/\mu^2 < (B/\mu^2)_I$, that map onto domain walls at zero temperature, constructed in section 6 (see figure 7). Note that the critical temperatures for the existence of the striped instabilities for the black holes is very low and we have not been able to stabilise the numerics to find their precise values. In section 7.2 we construct the superfluid instabilities both for the dyonic $AdS_2 \times \mathbb{R}^2$ solutions and for the thermodynamically preferred unbroken phase black holes in region I and we find that they only exist for $0 \leq B/\mu^2 \leq (B/\mu^2)_c < (B/\mu^2)_I$. In section 7.3 we discuss the implications for the full phase diagram of the system, which we summarised in figure 1.

7.1 Striped instabilities

We consider spatially modulated, or “striped”, perturbations about the dyonic $AdS_2 \times \mathbb{R}^2$ solutions given in section 3. We write these solutions as

$$\begin{aligned}
 ds^2 &= L^2 \left(-\rho^2 dt^2 + \rho^{-2} d\rho^2 \right) + dx_1^2 + dx_2^2, \\
 A &= -EL^2 \rho dt - \frac{1}{2} B x_2 dx_1 + \frac{1}{2} B x_1 dx_2, \quad \sigma = \sigma_0
 \end{aligned}
 \tag{7.1}$$

with σ_0 , E , L and B constants satisfying (3.2), (3.3). We consider the linearised perturbation

$$\begin{aligned}
 \delta g_{tt}(\rho, x_1) &= L^2 \rho^2 h_{tt}(\rho) \cos(kx_1), \\
 \delta g_{tx_2}(\rho, x_1) &= L^2 \rho h_{tx_2}(\rho) \sin(kx_1), \\
 \delta g_{x_1 x_1}(\rho, x_1) &= h_{x_1 x_1}(\rho) \cos(kx_1), \\
 \delta g_{x_2 x_2}(\rho, x_1) &= h_{x_2 x_2}(\rho) \cos(kx_1), \\
 \delta A_t(\rho, x_1) &= \rho \delta a_t(\rho) \cos(kx_1), \\
 \delta A_{x_2}(\rho, x_1) &= \delta a_{x_2}(\rho) \sin(kx_1), \\
 \delta \sigma(\rho, x_1) &= s(\rho) \cos(kx_1),
 \end{aligned}
 \tag{7.2}$$

where the wave-number k is a constant. Substituting into the equations of motion we are lead to a system of coupled linear differential equations that are second order in $h_{tt}(\rho)$, $h_{tx_2}(\rho)$, $a_t(\rho)$, $a_{x_2}(\rho)$, $s(\rho)$ and, when $k \neq 0$, first order in $h_{x_1 x_1}(\rho)$, $h_{x_2 x_2}(\rho)$. When $k = 0$ they are first order in $h_{x_2 x_2}(\rho)$ but second order in $h_{x_1 x_1}(\rho)$.

We are interested in analysing the spectrum of scaling dimensions of the operators in the CFT dual to the AdS_2 solution and in particular whether there are any modes that violate the BF bound. We therefore look for solutions where the seven functions of ρ are of the form $\mathbf{v} \rho^{-\delta}$ where \mathbf{v} is a constant vector and δ is a constant related to the scaling dimension via $\Delta = \delta$ or $\Delta = 1 - \delta$. The BF bound is given by $\Delta = 1/2$. The system of equations then takes the form $\mathbf{M} \mathbf{v} = 0$ where \mathbf{M} is a 7×7 matrix. Demanding that non-trivial \mathbf{v} exists implies that $\det \mathbf{M} = 0$ and this specifies the possible values of δ as a function k .

We first focus on the electric family of $AdS_2 \times \mathbb{R}^2$ solutions. We further restrict to the purely electric solution with $\sigma_0 = B = 0$. We find that the spectrum of δ is symmetric with respect to $k \rightarrow -k$ and we have displayed the values of Δ in figure 9 for $k \geq 0$. In particular, we find that, for $k \geq 0$, the BF bound is violated in the approximate range $k \in (5.94, 6.48)$, as already observed in [8]. We next analyse what happens as we move along the electric family of dyonic solutions by allowing $\sigma_0 \neq 0$. We find the surprising result that after rescaling $k \rightarrow \sqrt{\cosh(\sigma_0/\sqrt{3})} k$ the spectrum is independent of σ_0 . It would be interesting to better understand the underlying origin of this result. Moving to the magnetic family of $AdS_2 \times \mathbb{R}^2$ solutions, we find that the spectrum is identical to the electric family, which is presumably a consequence of the duality symmetry (2.4).

We conclude that all the dyonic $AdS_2 \times \mathbb{R}^2$ solutions of section 3 have spatially modulated instabilities.

7.2 Superfluid instabilities

To study superfluid instabilities we now consider the consistent Kaluza-Klein truncation of [24] which also includes the charged scalar field, denoted by χ in [24]. Writing $\chi =$

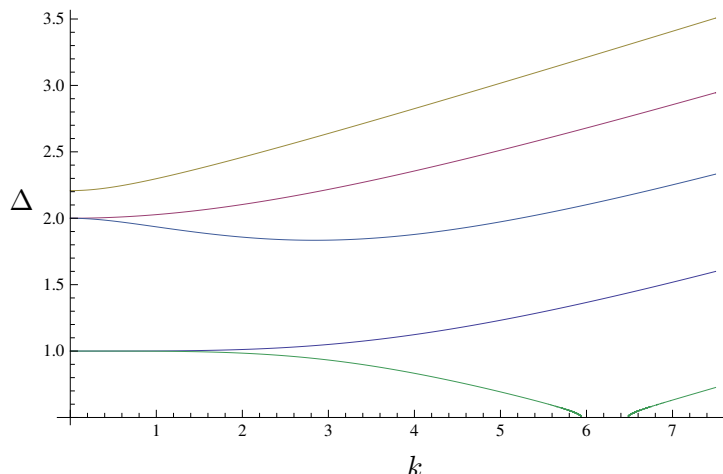


Figure 9. The spectrum of scaling dimensions Δ for spatially modulated striped modes of the purely electric $AdS_2 \times \mathbb{R}^2$ solution as a function of wave-number k . For $k \in (5.94, 6.48)$ the BF bound is violated and there are striped instabilities. After rescaling the momentum we find exactly the same spectrum for the entire family of dyonic $AdS_2 \times \mathbb{R}^2$ solutions.

$\sqrt{\frac{4}{3}} \tanh(\eta/2) e^{2i\theta}$, where η, θ are real, the action of [24] can be written

$$S = \frac{1}{16\pi G} \int d^4x \sqrt{-g} \left(R - \frac{1}{2} (\partial\sigma)^2 - \frac{\tau(\sigma)}{4} F^2 - V(\sigma, \eta) - \frac{1}{2} (\partial\eta)^2 - 2 \sinh^2 \eta (\partial\theta - A)^2 \right) + \frac{1}{32\pi G} \int \vartheta(\sigma) F \wedge F, \quad (7.3)$$

where $\tau(\sigma) = \frac{1}{\cosh \sqrt{3}\sigma}$ and $\vartheta(\sigma) = \tanh \sqrt{3}\sigma$ as before, and now

$$V(\sigma, \eta) \equiv -24 \cosh\left(\frac{\sigma}{\sqrt{3}}\right) \cosh^4\left(\frac{\eta}{2}\right) \left[1 - \frac{4}{3} \tanh^2\left(\frac{\eta}{2}\right) \cosh^2\left(\frac{\sigma}{\sqrt{3}}\right) \right]. \quad (7.4)$$

We want to consider linearised fluctuations of the charged scalar field about the dyonic solutions with $\eta = 0$, that we have constructed earlier. Similar computations were first carried out for other models in [3, 41]. We can consistently work in a gauge with $\theta = 0$. The linearised equation of motion for η then reads

$$[\nabla^2 - 4A^2 - f(\sigma)]\eta = 0, \quad (7.5)$$

where ∇ and A refer to the background solution and

$$f(\sigma) = 4 \cosh(\sqrt{3}\sigma) - 12 \cosh \frac{\sigma}{\sqrt{3}}. \quad (7.6)$$

First consider perturbations about the dyonic $AdS_2 \times \mathbb{R}^2$ solutions as given in (7.1). We focus on the lowest Landau level by writing

$$\eta(t, \rho, x_i) = e^{-\frac{|B|}{2}(x_1^2 + x_2^2)} \bar{\eta}(t, \rho). \quad (7.7)$$

We then find that $\bar{\eta}$ satisfies

$$[\nabla_{AdS_2}^2 - L^2 M^2] \bar{\eta} = 0, \quad M^2 = 2|B| + f(\sigma_0) - 4E^2 L^2, \quad (7.8)$$

where $\nabla_{AdS_2}^2$ is the Laplacian for a unit radius AdS_2 . This mode violates the AdS_2 BF bound when $L^2 M^2 < -1/4$. Focussing on $\sigma_0 \geq 0$, for the electric family we find that the mode is unstable apart from the range of solutions where, approximately, $|\sigma_0| \in (1.01, 2.94)$. We also find that for the magnetic family of solutions this mode violates the BF bound for (approximately) $|\sigma_0| \geq 3.08$. In section 5 we constructed domain walls for UV data μ, B which map onto domain walls that approach dyonic $AdS_2 \times \mathbb{R}^2$ solutions on either the electric or magnetic family in the IR. In figures 3 and 4 we have indicated which of these domain wall solutions must be unstable.

We can also consider linearised perturbations about the dyonic black hole solutions lying in the ansatz (4.1) that we constructed in section 6. In particular, we want to consider static, normalisable modes which appear at a critical temperature at which the superfluid instability appears. At this critical temperature new superfluid black hole solutions will appear. The mode that will have the highest critical temperature has the form

$$\eta = e^{-\frac{|B|}{2}(x_1^2 + x_2^2)} R(r), \quad (7.9)$$

where $R(r)$ satisfies the ODE

$$r^{-2} e^{\beta/2} \left(r^2 e^{-\beta/2} g R' \right)' - \left(\frac{2|B|}{r^2} + f(\sigma) - 4\phi^2 e^\beta g^{-1} \right) R = 0. \quad (7.10)$$

At the black hole event horizon we have that $R(r_h)$ is a constant and we can use the linearity of the ODE to set this to unity. At the AdS_4 boundary we demand that $R(r) = \eta_1/r + \eta_2/r^2 + \dots$ with $\eta_1 = 0$, corresponding to a spontaneous breaking of the abelian symmetry. We thus have fixed two integration constants for the second order ODE and hence we expect that solutions will appear at specific temperatures.

We are most interested in superfluid instabilities that appear on the thermodynamically preferred dyonic black hole solutions that we summarised in figure 7. The temperature at which the normalisable mode appears as a function of B/μ^2 are presented in figure 10 and we see that they exist in the range $0 \leq B/\mu^2 \leq (B/\mu^2)_c$ with $(B/\mu^2)_c < (B/\mu^2)_I$. We expect that $(B/\mu^2)_c$ is the same limiting value at which the thermodynamically preferred domain wall solutions on the electric branch have an instability in the $AdS_2 \times \mathbb{R}^2$ region in the IR; the latter is marked with a red dot in figure 10.

7.3 Conclusions about the full phase diagram

In region I, with $0 \leq B/\mu^2 < (B/\mu^2)_I$, the thermodynamically preferred unbroken phase black holes approach a domain wall solution at zero temperature that maps onto a dyonic $AdS_2 \times \mathbb{R}^2$ solution in the electric family at $T = 0$, as in figure 7. More precisely, they are domain walls that are on the upper part of the curve in the left hand plot in figure 3. For $0 \leq B/\mu^2 \leq (B/\mu^2)_c < (B/\mu^2)_I$ there are both superfluid and striped instabilities with the critical temperature for the superconducting instability monotonically decreasing to

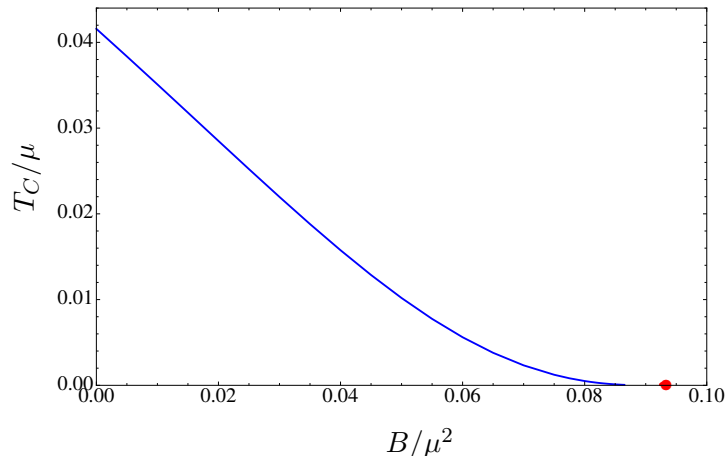


Figure 10. A plot of the critical temperature at which the superfluid phase transition sets in, for the thermodynamically preferred normal phase black holes summarised in figure 7 with $B/\mu^2 < (B/\mu^2)_I$. The red dot corresponds to the red dot appearing on the upper branch in the left hand plot of figure 3.

zero at $B/\mu^2 = (B/\mu^2)_c$. For $(B/\mu^2)_c \leq B/\mu^2 \leq (B/\mu^2)_I$ there are no longer superfluid instabilities but the striped instabilities persist. The critical temperature for the striped instabilities is very low for all values of B/μ^2 .

In order to deduce the phase diagram one needs to construct the fully back reacted black hole solutions for both the superfluid and the striped black holes. While for $B = 0$ the back-reacted superfluid black holes were constructed by solving ODEs [4], for $B \neq 0$ we need to solve PDEs.¹⁰ Similarly, constructing the back reacted striped black holes will also require solving PDEs, and this will be especially challenging because of the low critical temperature at which they appear.

To proceed, we therefore make some reasonable simplifying assumptions. Firstly, that all of the superfluid and striped black holes arise as second order transitions from the branch of unbroken phase black holes and secondly that they then continue down to $T = 0$ without sprouting additional branches. Finally, we also assume that the free energy curves only cross at most once as one lowers the temperature. To deduce the phase diagram of figure 1 then just requires a little thought concerning the free energy of the black holes.

For small B/μ^2 we stay on the superfluid branch all the way down to $T = 0$ as shown in figure 11 (a) and also in figure 1. For $(B/\mu^2)_c < B/\mu^2 < (B/\mu^2)_I$ there is no longer a branch of superfluid black holes and hence we must have a striped phase down to zero temperature, as depicted in figure 1. What happens near $B/\mu^2 = (B/\mu^2)_c$? One possibility is depicted in figure 11 (b): one first moves onto the superfluid phase via a second order transition and then onto the striped phase via a first order transition. This possibility is adopted in the phase diagram of figure 1 along with the tri-critical point that must necessarily appear and the first order transition at $T = 0$ at $B/\mu^2 = (B/\mu^2)_{(i)}$. A slightly

¹⁰It is likely that they form some kind of a vortex lattice e.g. [42]. One can speculate that at least for small B at $T = 0$ they become domain walls interpolating between two AdS_4 spaces, in order to match what happens at $B = 0$ [4].

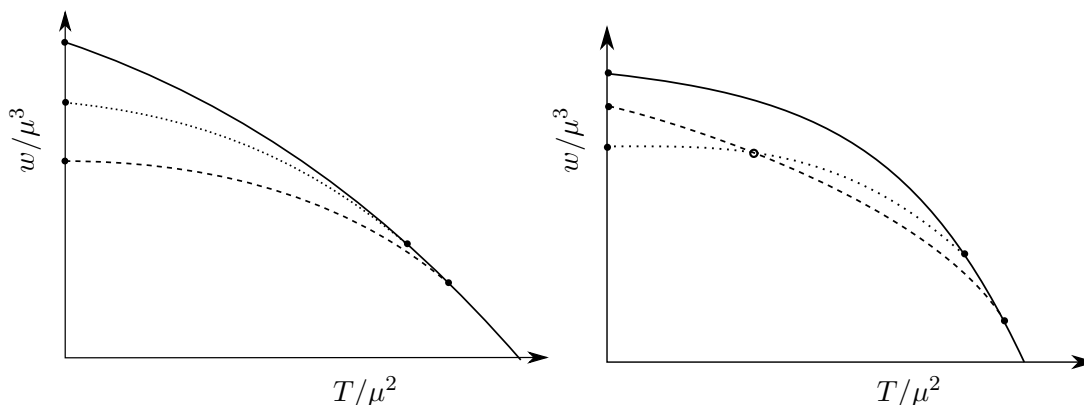


Figure 11. Schematic behaviour of the free energy versus temperature plots for the unbroken phase black holes (solid lines), the superfluid black holes (dashed lines) and the striped black holes (dotted lines). Figure (a) corresponds to values of B/μ^2 with $0 \leq B/\mu^2 \leq (B/\mu^2)_{(i)}$ in figure 1. Figure (b) corresponds to B/μ^2 being slightly bigger than $(B/\mu^2)_{(i)}$ in figure 1. One could also have the slightly alternative scenario where the roles of the superfluid and striped branches are interchanged in the crossover.

different possibility is that in figure 1 $(B/\mu^2)_{(i)}$ is closer to $(B/\mu^2)_c$ so that one would move from a striped phase via a first-order transition into a superfluid phase. At the interface of stripes and metamagnetism a very similar kind of reasoning, which we won't spell out implies that the striped black holes will exist for small temperatures up to $(B/\mu^2)_{(ii)}$ with $(B/\mu^2)_{(ii)} > (B/\mu^2)_I$ and generically $(B/\mu^2)_{(ii)} < (B/\mu^2)_{\max}$.

8 Final comments

We have shown that the phase structure of the $d = 3$ CFTs, as a function of T, μ, B , exhibit a rich phenomenology, summarised in figure 1. It would be very interesting to identify the $T = 0$ ground states that emerge in the superfluid and striped phases with $B \neq 0$. However, this appears to be a very challenging problem at the technical level. On the other hand we have successfully identified the $T \rightarrow 0$ hyperscaling violating ground states, with $z = 3/2, \theta = -2$, for larger values of the magnetic field. It would be desirable to directly construct the $T = 0$ domain wall solutions interpolating between the hyperscaling violating solutions (6.1) in the IR and AdS_4 in the UV.

All of the black hole solutions in this paper are not deformed in the UV by the operator \mathcal{O}_σ of dimension $\Delta = 2$, dual to the pseudo-scalar σ . Recall that the role of such deformations on the superfluid instability when $B = 0$ were analysed in [24], where a characteristic superfluid dome was seen. Outside of the dome was a class of charged black holes that as $T \rightarrow 0$ become singular in the IR. We have checked that this singularity actually corresponds to a hyperscaling violating behaviour with $z = 1$ and $\theta = -1$.

It will be very interesting to unify the analysis of this paper with that of [24] and study the $d = 3$ CFTs as a function of T, μ, B and deformation by \mathcal{O}_σ . It will be particularly interesting to see how the hyperscaling violating behaviours with $z = 3/2$ and

$z = 1$ interpolate between each other. It will also be very interesting to see if the deformation by \mathcal{O}_σ can drive down the temperature of the critical metamagnetic point to zero temperature to obtain a metamagnetic quantum critical point as observed, for example, in $Sr_3Ru_2O_7$ [43, 44] and studied in a different in a different holographic context in [45], building on [46, 47].

Acknowledgments

We would like to thank Alejandra Castro for interesting conversations. AD is supported by an EPSRC Postdoctoral Fellowship, JPG is supported by a Royal Society Wolfson Award and BW is supported by a Royal Commission for the Exhibition of 1851 Science Research Fellowship. This work was supported in part by STFC grant ST/J0003533/1. This work was supported in part by the U.S. Department of Energy (DOE) under cooperative research agreement Contract Number DE-FG02-05ER41360. BW thanks the CTP at MIT for hospitality while this paper was being finished.

A Dyonic AdS-RN black holes of Einstein-Maxwell theory

As somewhat of an aside, in this appendix we calculate the magnetisation and magnetic susceptibility of the canonical dyonic AdS-RN black holes of Einstein-Maxwell theory following [3, 38]. In particular, we want to contrast this with our results for the dyonic black holes in the top-down model (2.1), (2.2). To obtain Einstein-Maxwell theory, in (2.1) we set $\sigma = 0$ and

$$V \equiv -24, \quad \tau \equiv 1, \quad \vartheta \equiv 0. \quad (\text{A.1})$$

The dyonic AdS-RN black hole solution is given by

$$g = 4r^2 - \left(4r_+^2 + \frac{\mu^2}{4} + \frac{B^2}{4r_+^2}\right) \frac{r_+}{r} + \left(\frac{\mu^2}{4} + \frac{B^2}{4r_+^2}\right) \frac{r_+^2}{r^2}, \quad \phi = \mu \left(1 - \frac{r_+}{r}\right). \quad (\text{A.2})$$

The temperature is given by

$$T = \frac{48r_+^4 - \mu^2 r_+^2 - B^2}{16\pi r_+^3}, \quad (\text{A.3})$$

and this provides us with an expression $r_+ = r_+(T, \mu, B)$. Calculating the free-energy $w = w(T, \mu, B)$ as in section 4.3 we obtain

$$w = -4r_+^3 - \frac{\mu^2 r_+}{4} + \frac{3B^2}{4r_+}, \quad (\text{A.4})$$

and for the magnetisation m and magnetic susceptibility χ_m at constant T, μ we get

$$\begin{aligned} m &= - \left(\frac{\partial w}{\partial B} \right)_{T, \mu} = - \frac{B}{r_+}, \\ \chi_m &= \left(\frac{\partial m}{\partial B} \right)_{T, \mu} = - \frac{48r_+^4 + \mu^2 r_+^2 + B^2}{r_+ (48r_+^4 + \mu^2 r_+^2 + 3B^2)}. \end{aligned} \quad (\text{A.5})$$

Notice that since m is always negative the system is strongly diamagnetic (as noted in [3]). Also, χ_m is always negative asymptoting to zero for large B, T .

We can also use the ensemble $f = f(T, q, B) \equiv w + \mu q$. We find

$$f = -4r_+^3 + \frac{3(q^2 + B^2)}{4r_+}, \tag{A.6}$$

and

$$m = - \left(\frac{\partial f}{\partial B} \right)_{T,q} = - \frac{B}{r_+},$$

$$\chi'_m = \left(\frac{\partial m}{\partial B} \right)_{T,q} = - \frac{48r_+^4 + 3q^2 + B^2}{r_+(48r_+^4 + 3(q^2 + B^2))}, \tag{A.7}$$

and we note that χ'_m has similar behaviour to χ_m , as one expects.

References

- [1] S.S. Gubser, *Breaking an Abelian gauge symmetry near a black hole horizon*, *Phys. Rev. D* **78** (2008) 065034 [[arXiv:0801.2977](#)] [[INSPIRE](#)].
- [2] S.A. Hartnoll, C.P. Herzog and G.T. Horowitz, *Building a Holographic Superconductor*, *Phys. Rev. Lett.* **101** (2008) 031601 [[arXiv:0803.3295](#)] [[INSPIRE](#)].
- [3] S.A. Hartnoll, C.P. Herzog and G.T. Horowitz, *Holographic Superconductors*, *JHEP* **12** (2008) 015 [[arXiv:0810.1563](#)] [[INSPIRE](#)].
- [4] J.P. Gauntlett, J. Sonner and T. Wiseman, *Holographic superconductivity in M-theory*, *Phys. Rev. Lett.* **103** (2009) 151601 [[arXiv:0907.3796](#)] [[INSPIRE](#)].
- [5] S.S. Gubser, C.P. Herzog, S.S. Pufu and T. Tesileanu, *Superconductors from Superstrings*, *Phys. Rev. Lett.* **103** (2009) 141601 [[arXiv:0907.3510](#)] [[INSPIRE](#)].
- [6] S. Nakamura, H. Ooguri and C.-S. Park, *Gravity Dual of Spatially Modulated Phase*, *Phys. Rev. D* **81** (2010) 044018 [[arXiv:0911.0679](#)] [[INSPIRE](#)].
- [7] H. Ooguri and C.-S. Park, *Holographic End-Point of Spatially Modulated Phase Transition*, *Phys. Rev. D* **82** (2010) 126001 [[arXiv:1007.3737](#)] [[INSPIRE](#)].
- [8] A. Donos and J.P. Gauntlett, *Holographic striped phases*, *JHEP* **08** (2011) 140 [[arXiv:1106.2004](#)] [[INSPIRE](#)].
- [9] A. Donos and J.P. Gauntlett, *Holographic helical superconductors*, *JHEP* **12** (2011) 091 [[arXiv:1109.3866](#)] [[INSPIRE](#)].
- [10] A. Donos, J.P. Gauntlett and C. Pantelidou, *Spatially modulated instabilities of magnetic black branes*, *JHEP* **01** (2012) 061 [[arXiv:1109.0471](#)] [[INSPIRE](#)].
- [11] A. Donos, J.P. Gauntlett and C. Pantelidou, *Magnetic and Electric AdS Solutions in String- and M-theory*, *Class. Quant. Grav.* **29** (2012) 194006 [[arXiv:1112.4195](#)] [[INSPIRE](#)].
- [12] S.K. Domokos and J.A. Harvey, *Baryon number-induced Chern-Simons couplings of vector and axial-vector mesons in holographic QCD*, *Phys. Rev. Lett.* **99** (2007) 141602 [[arXiv:0704.1604](#)] [[INSPIRE](#)].
- [13] H. Ooguri and C.-S. Park, *Spatially Modulated Phase in Holographic quark-gluon Plasma*, *Phys. Rev. Lett.* **106** (2011) 061601 [[arXiv:1011.4144](#)] [[INSPIRE](#)].

- [14] C.B. Bayona, K. Peeters and M. Zamaklar, *A Non-homogeneous ground state of the low-temperature Sakai-Sugimoto model*, *JHEP* **06** (2011) 092 [[arXiv:1104.2291](#)] [[INSPIRE](#)].
- [15] O. Bergman, N. Jokela, G. Lifschytz and M. Lippert, *Striped instability of a holographic Fermi-like liquid*, *JHEP* **10** (2011) 034 [[arXiv:1106.3883](#)] [[INSPIRE](#)].
- [16] A. Donos and J.P. Gauntlett, *Helical superconducting black holes*, *Phys. Rev. Lett.* **108** (2012) 211601 [[arXiv:1203.0533](#)] [[INSPIRE](#)].
- [17] A. Donos and J.P. Gauntlett, *Black holes dual to helical current phases*, *Phys. Rev. D* **86** (2012) 064010 [[arXiv:1204.1734](#)] [[INSPIRE](#)].
- [18] N. Iizuka et al., *Bianchi Attractors: a Classification of Extremal Black Brane Geometries*, *JHEP* **07** (2012) 193 [[arXiv:1201.4861](#)] [[INSPIRE](#)].
- [19] M. Rozali, D. Smyth, E. Sorkin and J.B. Stang, *Holographic Stripes*, [arXiv:1211.5600](#) [[INSPIRE](#)].
- [20] G. Lifschytz and M. Lippert, *Holographic Magnetic Phase Transition*, *Phys. Rev. D* **80** (2009) 066007 [[arXiv:0906.3892](#)] [[INSPIRE](#)].
- [21] C. Charmousis, B. Gouteraux, B. Kim, E. Kiritsis and R. Meyer, *Effective Holographic Theories for low-temperature condensed matter systems*, *JHEP* **11** (2010) 151 [[arXiv:1005.4690](#)] [[INSPIRE](#)].
- [22] N. Ogawa, T. Takayanagi and T. Ugajin, *Holographic Fermi Surfaces and Entanglement Entropy*, *JHEP* **01** (2012) 125 [[arXiv:1111.1023](#)] [[INSPIRE](#)].
- [23] L. Huijse, S. Sachdev and B. Swingle, *Hidden Fermi surfaces in compressible states of gauge-gravity duality*, *Phys. Rev. B* **85** (2012) 035121 [[arXiv:1112.0573](#)] [[INSPIRE](#)].
- [24] J.P. Gauntlett, J. Sonner and T. Wiseman, *Quantum Criticality and Holographic Superconductors in M-theory*, *JHEP* **02** (2010) 060 [[arXiv:0912.0512](#)] [[INSPIRE](#)].
- [25] X. Dong, S. Harrison, S. Kachru, G. Torroba and H. Wang, *Aspects of holography for theories with hyperscaling violation*, *JHEP* **06** (2012) 041 [[arXiv:1201.1905](#)] [[INSPIRE](#)].
- [26] H. Singh, *Lifshitz/Schrödinger Dp-branes and dynamical exponents*, *JHEP* **07** (2012) 082 [[arXiv:1202.6533](#)] [[INSPIRE](#)].
- [27] K. Narayan, *On Lifshitz scaling and hyperscaling violation in string theory*, *Phys. Rev. D* **85** (2012) 106006 [[arXiv:1202.5935](#)] [[INSPIRE](#)].
- [28] P. Dey and S. Roy, *Lifshitz-like space-time from intersecting branes in string/M theory*, *JHEP* **06** (2012) 129 [[arXiv:1203.5381](#)] [[INSPIRE](#)].
- [29] J.P. Gauntlett, S. Kim, O. Varela and D. Waldram, *Consistent supersymmetric Kaluza-Klein truncations with massive modes*, *JHEP* **04** (2009) 102 [[arXiv:0901.0676](#)] [[INSPIRE](#)].
- [30] M. Duff, B. Nilsson and C. Pope, *The criterion for vacuum stability in Kaluza-Klein supergravity*, *Phys. Lett. B* **139** (1984) 154 [[INSPIRE](#)].
- [31] A. Imaanpur and M. Naghdi, *Dual Instantons in Anti-membranes Theory*, *Phys. Rev. D* **83** (2011) 085025 [[arXiv:1012.2554](#)] [[INSPIRE](#)].
- [32] D. Forcella and A. Zaffaroni, *Non-supersymmetric CS-matter theories with known AdS duals*, *Adv. High Energy Phys.* **2011** (2011) 393645 [[arXiv:1103.0648](#)] [[INSPIRE](#)].
- [33] F. Denef and S.A. Hartnoll, *Landscape of superconducting membranes*, *Phys. Rev. D* **79** (2009) 126008 [[arXiv:0901.1160](#)] [[INSPIRE](#)].

- [34] M. Taylor, *Non-relativistic holography*, [arXiv:0812.0530](#) [[INSPIRE](#)].
- [35] S.S. Gubser and F.D. Rocha, *Peculiar properties of a charged dilatonic black hole in AdS₅*, *Phys. Rev. D* **81** (2010) 046001 [[arXiv:0911.2898](#)] [[INSPIRE](#)].
- [36] K. Goldstein, S. Kachru, S. Prakash and S.P. Trivedi, *Holography of Charged Dilaton Black Holes*, *JHEP* **08** (2010) 078 [[arXiv:0911.3586](#)] [[INSPIRE](#)].
- [37] K. Goldstein et al., *Holography of Dyonically Charged Dilaton Black Branes*, *JHEP* **10** (2010) 027 [[arXiv:1007.2490](#)] [[INSPIRE](#)].
- [38] F. Denef, S.A. Hartnoll and S. Sachdev, *Quantum oscillations and black hole ringing*, *Phys. Rev. D* **80** (2009) 126016 [[arXiv:0908.1788](#)] [[INSPIRE](#)].
- [39] S.A. Hartnoll and P. Kovtun, *Hall conductivity from dyonic black holes*, *Phys. Rev. D* **76** (2007) 066001 [[arXiv:0704.1160](#)] [[INSPIRE](#)].
- [40] V. Balasubramanian and P. Kraus, *A Stress tensor for Anti-de Sitter gravity*, *Commun. Math. Phys.* **208** (1999) 413 [[hep-th/9902121](#)] [[INSPIRE](#)].
- [41] T. Albash and C.V. Johnson, *A Holographic Superconductor in an External Magnetic Field*, *JHEP* **09** (2008) 121 [[arXiv:0804.3466](#)] [[INSPIRE](#)].
- [42] K. Maeda, M. Natsuume and T. Okamura, *Vortex lattice for a holographic superconductor*, *Phys. Rev. D* **81** (2010) 026002 [[arXiv:0910.4475](#)] [[INSPIRE](#)].
- [43] S.A. Grigera et al., *Disorder-Sensitive Phase Formation Linked to Metamagnetic Quantum Criticality*, *Science* **12** (2004) 1154.
- [44] A.J. Millis, A.J. Schofield, G.G. Lonzarich and S.A. Grigera, *Metamagnetic Quantum Criticality in Metals*, *Physical Review Letters* **88** (2002) 217204 [[cond-mat/0109440](#)].
- [45] E. D'Hoker and P. Kraus, *Holographic Metamagnetism, Quantum Criticality and Crossover Behavior*, *JHEP* **05** (2010) 083 [[arXiv:1003.1302](#)] [[INSPIRE](#)].
- [46] E. D'Hoker and P. Kraus, *Magnetic Brane Solutions in AdS*, *JHEP* **10** (2009) 088 [[arXiv:0908.3875](#)] [[INSPIRE](#)].
- [47] E. D'Hoker and P. Kraus, *Charged Magnetic Brane Solutions in AdS₅ and the fate of the third law of thermodynamics*, *JHEP* **03** (2010) 095 [[arXiv:0911.4518](#)] [[INSPIRE](#)].