

# Signature Change in Loop Quantum Cosmology

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**Abstract** The Wick rotation is commonly considered only as a useful computational trick. However, as suggested by Hartle and Hawking already in early eighties, Wick rotation may gain physical meaning at the Planck epoch. While such possibility is conceptually interesting, leading to no-boundary proposal, mechanism behind the signature change remains mysterious. We show that the signature change anticipated by Hartle and Hawking naturally appears in loop quantum cosmology. Theory of cosmological perturbations with the effects of quantum holonomies is discussed. It was shown by Cailleteau et al. [3] that this theory can be uniquely formulated in an anomaly-free manner. The obtained algebra of effective constraints turns out to be modified so that the metric signature is changing from Lorentzian in low curvature regime to Euclidean in high curvature regime. Implications of this phenomenon on propagation of cosmological perturbations are discussed and corrections to inflationary power spectra of scalar and tensor perturbations are derived. Possible relations with other approaches to quantum gravity are outlined. We also propose an intuitive explanation of the observed signature change using analogy with spontaneous symmetry breaking in “wired” metamaterials.

## 1 Introduction

The metric signature change from Lorentzian to Euclidean is usually performed by the so-called Wick rotation ( $t \rightarrow -i\tau$ ), under which the line element  $ds^2 = -dt^2 + d\mathbf{x}^2$  transforms to  $ds^2 = d\tau^2 + d\mathbf{x}^2$ . The Wick rotation becomes especially important in the path integral formulation of quantum mechanics. It allows to calculate non-perturbative effects by considering *instantons*. Another advantage coming from the

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Wick rotation is improvement of a convergence property of some path integrals. But these are just useful computational tricks.

However, in 1983 Hartle and Hawking proposed that Wick rotation may gain physical meaning at the Planck epoch [1]. This assumption was crucial for construction of the so-called no-boundary proposal, which was a way to cope with the problem of initial conditions for the Universe. While such possibility is conceptually interesting, mechanism behind the signature change in the very early Universe remains enigmatic.

If such transition from the Lorentzian to Euclidean space has occurred in the early Universe, what could be the origin of this? Can the signature change be due to some quantum gravity effects? So far, there have been no indications supporting such possibility. However, recent results coming from symmetry reduced models of Loop Quantum Gravity (LQG) [2] suggest that indeed the signature change may occur due to the discrete nature of space at the Planck scale [3, 4].

## 2 Loops

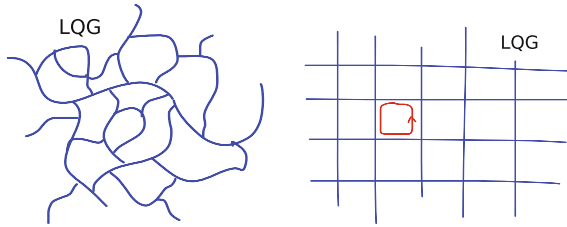
In LQG, granularity of space at the Planck scale is manifested by discrete spectra of geometric operators such as area and volume. The starting point for constructing LQG is the Hamiltonian formulation of General Relativity in language of Ashtekar variables fulfilling  $\{E_j^a(\mathbf{x}), A_b^i(\mathbf{y})\} = 8\pi G\gamma\delta_b^a\delta_j^i\delta^{(3)}(\mathbf{x} - \mathbf{y})$ , where  $\gamma$  is a free parameter of the theory called Barbero-Immirzi parameter. In this framework, Hamiltonian of gravity sector can be written as a sum of three constraints:  $H_G[N, N^a, N^i] = S[N] + D[N^a] + G[N^i] \approx 0$ . Here  $S$  is the scalar constraint,  $D$  is the diffeomorphism constraint and  $G$  is the Gauss constraint. The constraints ( $S \rightarrow \mathcal{C}_1, D \rightarrow \mathcal{C}_2, G \rightarrow \mathcal{C}_3$ ) fulfill the closed algebra  $\{\mathcal{C}_I, \mathcal{C}_J\} = f^K{}_{IJ}(A_b^j, E_i^a)\mathcal{C}_K$ , where  $f^K{}_{IJ}(A_b^j, E_i^a)$  are some structure functions.

Based on the Ashtekar variables, nonlocal variables called holonomies and fluxes are constructed. These new variables are the subject of quantization in LQG. For our purposes, it is sufficient to note that holonomy is defined as parallel transport of  $A_a = A_a^i\tau_i$  along some curve  $e$  on a spatial hypersurface:  $h_e := \mathcal{P} \exp \int_e A_a dx^a$ , where  $\sigma_j = 2i\tau_j$  are Pauli matrices. The holonomies are elements of  $SU(2)$  group.

In LQG, a state of gravity is described by superposition of graphs called *spin networks*. The links of the graphs are labelled by half integers ( $j = 1/2, 1, 3/2, \dots$ ) corresponding to irreducible representations of the  $SU(2)$  group. An exemplary spin network is shown in Fig. 1.

Loop Quantum Cosmology (LQC) [5] is a regular lattice model of LQG. In particular, symmetries of isotropy and homogeneity are imposed on the spin network within LQC. In what follows, we will consider an isotropic model with small perturbative inhomogeneities around a flat Friedmann-Robertson-Walker (FRW) background.

In LQC, physical area of the elementary lattice cell  $\text{Ar}_\square = \bar{p}\bar{\mu}^2$ , where  $\bar{p} = a^2$  and  $a$  is a scale factor. In general  $\bar{\mu} \propto \bar{p}^\beta$ , where  $-1/2 \leq \beta \leq 0$ . For the so-called



**Fig. 1** In LQG, state of geometry is described by the spin network (*left*). In LQC, the spin network takes the form of regular lattice (*right*). The (*red*) loop represents holonomy  $h_{\square}$  around the elementary cell

$\bar{\mu}$ -scheme (“new quantization scheme”):  $\bar{\mu} = \sqrt{\frac{\Delta}{\bar{p}}}$ , where usually  $\Delta$  is assumed to be area gap derived from LQG:  $\Delta = 2\sqrt{3}\pi\gamma l_{Pl}^2$ . The  $\bar{\mu}$ -scheme, in which physical area of the elementary lattice cell is constant during the cosmological evolution, was shown to be physically favoured.

At the effective level, effects of discreteness can be studied by introducing the so-called holonomy corrections. They are obtained by replacing curvature of Ashtekar connection by holonomy around elementary loop (see Fig. 1). This procedure is called *polymerization*.

### 3 Cosmological Perturbations

In most cosmological applications, Ashtekar variables can be decomposed into the background (here flat FRW) and perturbation parts:  $E_i^a = \bar{E}_i^a + \delta E_i^a$  and  $A_a^i = \bar{A}_a^i + \delta A_a^i$ , where  $\bar{E}_i^a = \bar{p}\delta_i^a$  and  $\bar{A}_a^i = \gamma\bar{k}\delta_a^i$ . The perturbations of the Ashtekar variables can be related with the standard metric perturbations: scalar modes ( $\Phi, \Psi, E, B$ ), vector modes ( $S^a, F_a$ ) and tensor modes ( $h_{ab}$ ). In total, there are 10 perturbative degrees of freedom. Furthermore, matter degrees of freedom are also subject of perturbative decomposition. In what follows we consider a model with a scalar field, so  $\varphi$  and its canonically conjugated momenta  $\pi$  can be written as:  $\varphi = \bar{\varphi} + \delta\varphi$  and  $\pi = \bar{\pi} + \delta\pi$ . Applying the above decompositions, total constraints  $\mathcal{C}_{tot} = \mathcal{C}_G + \mathcal{C}_M$ , which take into account contributions from gravity and matter, can be expanded. Our analysis is performed up to the second order in perturbative development:  $\mathcal{C}_{tot} = \mathcal{C}^{(0)} + \mathcal{C}^{(1)} + \mathcal{C}^{(2)} + \dots$ , so the corresponding equations of motion stay linear.

### 4 Anomaly Freedom and Algebra of Constraints

The effects of discreteness of space are introduced by employing the holonomy corrections. Such corrections modify the classical constraints  $\mathcal{C}_{tot}$  to some new effective quantum constraints  $\mathcal{C}_{tot}^Q$ . The modification follows the correspondence principle, such that in the limit  $\bar{\mu} \rightarrow 0$ , the modified constraints  $\mathcal{C}_{tot}^Q \rightarrow \mathcal{C}_{tot}$ . The procedure of introducing quantum corrections suffers from various ambiguities. Moreover, the resulting algebra of modified constraints is in general not closed:

$$\{\mathcal{C}_I^Q, \mathcal{C}_J^Q\} = g^{KIJ}(A_b^j, E_i^a)\mathcal{C}_K^Q + \mathcal{A}_{IJ}, \tag{1}$$

where  $\mathcal{A}_{IJ}$  are some anomaly terms. Closure of algebra is required by mathematical consistency of the theory. So, the question is: Can we introduce quantum holonomy corrections in the anomaly-free manner (i.e. such that  $\mathcal{A}_{IJ} = 0$ )? The answer turns out to be “yes”. Moreover, there is a unique way of modifying constraints such that the algebra is closed [3]. Additionally, the conditions of anomaly-freedom are fulfilled if and only if  $\beta = -1/2$ , which corresponds to the  $\bar{\mu}$ -scheme. Therefore, the only remaining free parameter is the area of elementary lattice cell  $\Delta$ , which is however expected to be of the order of the Planck area  $l_{Pl}^2$ .

The obtained algebra of the effective quantum constraint is [3]:

$$\{D_{tot}[N_1^a], D_{tot}[N_2^a]\} = 0, \tag{2}$$

$$\{S_{tot}[N], D_{tot}[N^a]\} = -S_{tot}[\delta N^a \partial_a \delta N], \tag{3}$$

$$\{S_{tot}[N_1], S_{tot}[N_2]\} = \Omega D_{tot} \left[ \frac{\bar{N}}{\bar{P}} \partial^a (\delta N_2 - \delta N_1) \right]. \tag{4}$$

The algebra is closed but deformed with respect to the classical case due to the presence of  $\Omega$  in (4). Therefore, general covariance is modified. The new factor  $\Omega$  can be expressed as follows:  $\Omega = \cos(2\bar{\mu}\gamma\bar{k}) = 1 - 2\rho/\rho_c \in [-1, 1]$ , where  $\rho$  is energy density of the scalar matter and the critical energy density  $\rho_c := \frac{3}{8\pi G\Delta} \sim \rho_{Pl} := m_{Pl}^4$ . What is the interpretation of the above deformation of the algebra of constraints? In order to answer this question let us recall the classical equivalent of the modified bracket (4) for a space with signature  $\sigma$  [2]:

$$\{S_{tot}[N_1], S_{tot}[N_2]\} = \sigma D \left[ \frac{\bar{N}}{\bar{P}} \partial^a (\delta N_2 - \delta N_1) \right].$$

Here,  $\sigma = 1$  corresponds to the Lorentzian signature and  $\sigma = -1$  to the Euclidean one. Therefore, we conclude that modification of the effective algebra of constraints (4) means that the space becomes Euclidean for  $\rho > \rho_c/2$ , while Lorentzian geometry emerges for  $\rho < \rho_c/2$ . In the regime of high curvatures and high energy densities ( $\rho > \rho_c/2$ ), spacetime becomes 4-dimensional Euclidean space. There is no distinguished time direction in this phase. It is interesting to notice that this model exhibits

properties of the Hartle-Hawking no-boundary proposal. However here signature change occurs smoothly with an increase of energy density.

The same effect of signature change was observed also for inhomogeneous spherically symmetric models with holonomy corrections [4]. Therefore, we have a good reason to believe that this phenomenon is a general consequence of quantum polymerization of space. Moreover, we can speculate that the *off-shell* algebra of quantum constraints in LQG should also exhibit such property, i.e.  $[\hat{S}, \hat{S}] = i\Omega \hat{D}$ .

Among many other comments, the fact that *ultralocal gravity* [6], where  $\{S, S\} = 0$ , is recovered at the transition point  $\rho = \rho_c/2$  is worth stressing.

## 5 Equations of Motion

The obtained anomaly-free formulation can be now used to derive equations of motion for both background variables and perturbations. The background dynamics is governed by the modified Friedmann equation  $H^2 = \frac{8\pi G}{3}\rho(1 - \rho/\rho_c)$ , where  $H$  is the Hubble factor. Clearly, only  $\rho \leq \rho_c$  are physically allowed, which was used to determine the range of  $\Omega$ .

For the scalar perturbations one can define a gauge-invariant variable  $v$  and the corresponding modified Mukhanov equation [3]:

$$\frac{d^2}{d\tau^2}v - \Omega \nabla^2 v - \frac{z''}{z}v = 0, \quad (5)$$

where  $z := \sqrt{\bar{p}} \frac{\dot{\phi}}{H}$ . Here,  $\tau = \int dt/a$  is a conformal time. For the considered model with a scalar field, vector modes are pure gauge, and therefore do not contribute [7]. Equation for tensor modes takes the form [8]

$$\frac{d^2}{d\tau^2}h_{ab} + 2 \left( aH - \frac{1}{2\Omega} \frac{d\Omega}{d\tau} \right) \frac{d}{d\tau}h_{ab} - \Omega \nabla^2 h_{ab} = 0.$$

The obtained equations are modified by the presence of  $\Omega$  in front of the Laplace operator. Therefore, transition to the Euclidean domain leads to a change of equation type from hyperbolic to elliptic, as expected. Furthermore, we see that the speed of propagation is varying, since  $c_s^2 = \Omega$ .

## 6 Holonomy Corrections to Inflationary Power Spectra

As an application of the obtained equations of motion we will derive holonomy corrections to the inflationary scalar and tensor power spectra. We will focus on the slow-roll inflationary model driven by a single scalar field  $\varphi$  with potential  $V(\varphi)$

occurring in the Lorentzian domain. The slow-roll parameters with the holonomy corrections are:

$$\varepsilon := \frac{m_{Pl}^2}{16\pi} \left( \frac{V_{,\varphi}}{V} \right)^2 \frac{1}{(1 - V/\rho_c)} \quad \text{and} \quad \eta := \frac{m_{Pl}^2}{8\pi} \left( \frac{V_{,\varphi\varphi}}{V} \right) \frac{1}{(1 - V/\rho_c)}.$$

Derivation of the scalar and tensor power spectra is based on application of the standard techniques of the quantum field theory in curved spaces. Moreover, normalization is such that in the UV limit the Minkowski vacuum is recovered. Obtained spectra of scalar and tensor (gravitational waves) perturbations are

$$\mathcal{P}_S(k) = A_S \left( \frac{k}{aH} \right)^{n_S-1} \quad \text{and} \quad \mathcal{P}_T(k) = A_T \left( \frac{k}{aH} \right)^{n_T},$$

where amplitudes and spectral indices are given as follows:

$$A_S = \frac{1}{\pi\varepsilon} \left( \frac{H}{m_{Pl}} \right)^2 \left( 1 + 2\frac{V}{\rho_c} \right) \quad \text{and} \quad n_S = 1 + 2\eta - 6\varepsilon \left( 1 - \frac{V}{\rho_c} \right),$$

$$A_T = \frac{16}{\pi} \left( \frac{H}{m_{Pl}} \right)^2 \left( 1 + 3\frac{V}{\rho_c} \right) \quad \text{and} \quad n_T = -2\varepsilon \left( 1 - 3\frac{V}{\rho_c} \right).$$

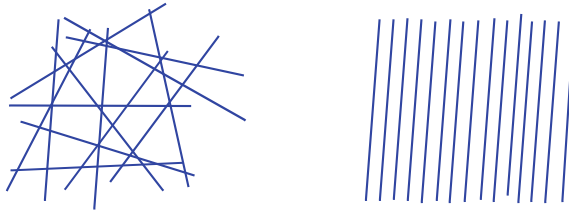
Furthermore, the consistency relation is

$$r := \frac{A_T}{A_S} \simeq 16\varepsilon \left( 1 + \frac{V}{\rho_c} \right). \quad (6)$$

The corrections are introduced by the factors  $V/\rho_c$ , which are of the order of  $10^{-12}$  for typical values of parameters. Confrontation of the obtained spectra with the available CMB data will be studied elsewhere [9]. In a more detailed analysis, initial conditions should be established at the transition point  $\rho = \rho_c/2$ , which can lead to some additional modification of the power spectra. This issue will be investigated in our further research.

## 7 Towards Understanding the Signature Change

It is tempting to understand the origin of the signature change at a microscopic level. Is this a kind of phase transition occurring at the level of the spin network or, more generally, at the level of some four dimensional spin foam model? One possibility is that some sort of spontaneous symmetry breaking takes place, leading to distinction of the time dimension in low curvature regime.



**Fig. 2** At high temperatures orientation of nanowires is random (*left*). At low temperatures, some direction is distinguished due to the spontaneous symmetry breaking (*right*). In the distinguished direction, dielectric permittivity is negative, leading to emergence of a time variable

A relevant example of spontaneous symmetry breaking is given by ferromagnets. At high temperatures, ferromagnets are losing their magnetic properties. The spins (magnetic moments) are randomly orientated, and no direction is distinguished. The system satisfies the  $SO(3)$  rotational symmetry, which is also a symmetry of the corresponding Hamiltonian. However, while temperature is decreased, spins start to orientate in some direction and magnetic domains are formed. This process begins when temperature is lowered below the so-called Curie temperature.

One can speculate that an analogous phase transition occurs in case of gravity. Namely, at high energies, the symmetry is, say,  $SO(4)$ . There is no distinction of any time coordinate. However, while lowering the energy density, which is an analogue of temperature, the symmetry will be broken to  $SO(3)$ . The energy density  $\rho_c/2$  is an analogue of the Curie temperature. It is possible that  $SO(3)$  symmetry of the triad rotations in Ashtekar formalism (doubly covered by  $SU(2)$  group in LQG) is in fact a residual symmetry of some wider symmetry before spontaneous symmetry breaking.

The phase transition, similar to the one discussed in case of ferromagnets, may occur for the so-called “wired” metamaterials composed of nanowires. It was shown that for such materials, an effective emergence of time variable may occur because of *negative dielectric permittivity* [10]. At high temperatures, dielectric permittivities in all directions are positive. However in a low temperature state, nanowires may align in some direction as spins do. In this distinguished direction, dielectric permittivity becomes negative leading to emergence of “time” direction at the level of equations of motion for electromagnetic field (see Fig. 2). We speculate that the same kind of process occurs in case of gravity, leading to emergence of a time variable. Structure of four dimensional Euclidean space undergoes phase transition, such that some particular direction is picked (presumably domains with different time directions can form). Equations of motion for the fields living on such frame change from elliptic to hyperbolic, which is interpreted as emergence of the time direction. We will explore this interpretation in more details in our further studies.

## 8 Summary and Outlook

We have shown that metric signature change may occur due to polymerization of space at the Planck scale. Preliminary analysis of this new phenomenon was carried out. Many questions remain open and are awaiting detailed analysis. In particular: Is there signature change in full LQG too? What are the initial conditions at  $\rho = \rho_c/2$ ? What is happening at the microscopic scale? How is the propagation of high energy photons affected? And many, many others.

Summarizing, the paradigm shift seems to be observed in LQC. There is no longer deterministic bouncing phase as was thought for many years. The Big Bounce model in LQC seems to be an artifact of the strong assumption of homogeneity. Due to the Euclidean stage there is no access to the information contained in contacting Universe, which gives answer to the long standing debate on cosmic forgetfulness.

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